LA-UR-22-21110

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Title: Next-Generation Simulations of The Remarkable Deaths of Massive Stars

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Intended for: oskar klein center online colloquium feb 8

Issued: 2022-02-09

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NEXT-GENERATION SIMULATIONS OF THE REMARKABLE DEATHS OF MASSIVE STARS

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OKC Colloquium
February 8th, 2022

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OVERVIEW

Introduction

- Core-Collapse Supernovae
- CCSN Explosion Mechanism
- The CCSN "Problem" and possible solutions

3D CCSN Progenitors

- Landscape of 3D Progenitors
- 3D Rotating 16 $M_{\odot}$ star
- Signals from CCSNe

Conclusions & Summary

RCW 114, an old supernova remnant with an estimated diameter of 100 lightyears.

INTRODUCTION

Core-Collapse Supernovae

CORE COLLAPSE SUPERNOVAE

Understanding core-collapse supernova explosions is crucial to many different problems of astronomy.

Galactic Chemical Evolution

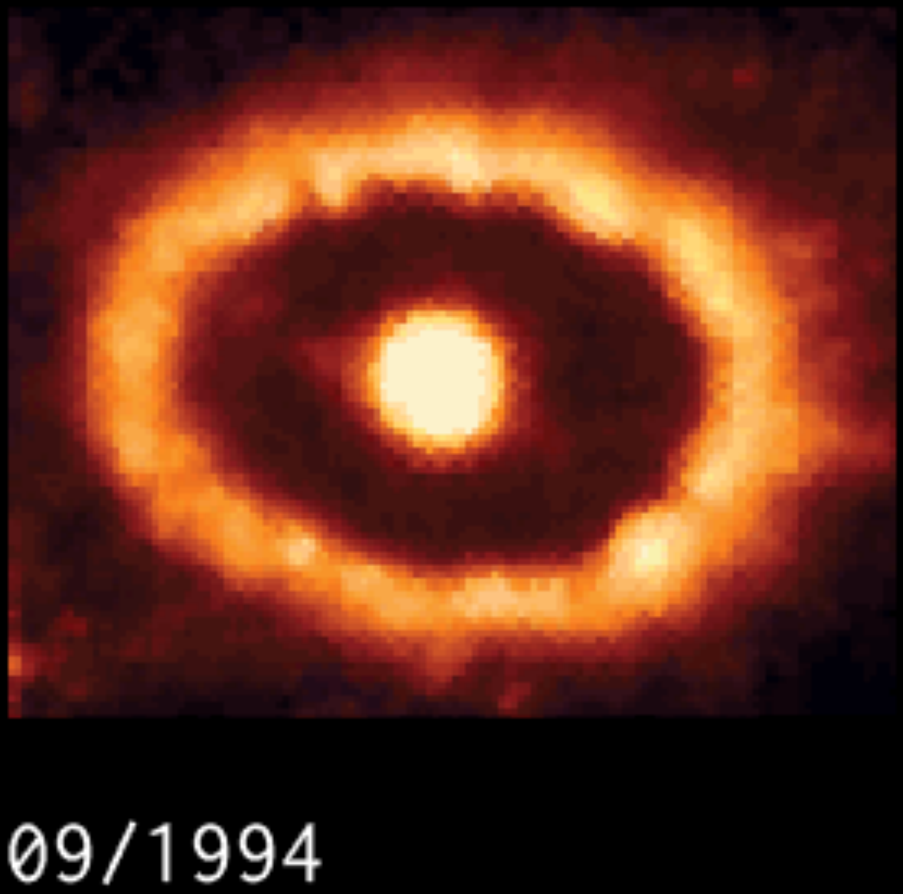
- Nucleosynthesis
- Stellar Feedback

Compact Object Formation

- Produce NS / stellar mass BHs

Multi-Messenger Astronomy

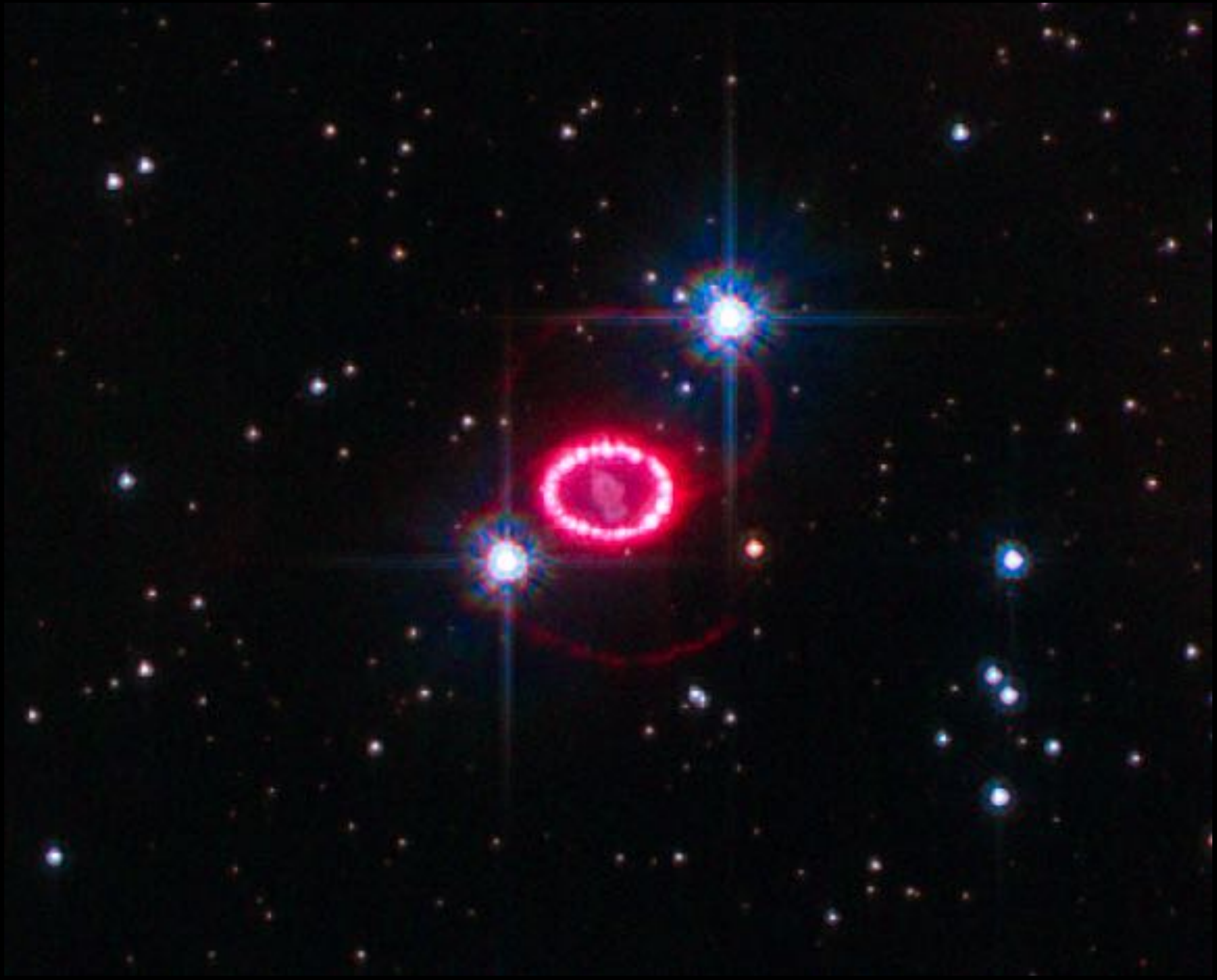
- Gravitational Waves
- Neutrino Emission



Credit: Larsson, J. et al. (2011).

CORE-COLLAPSE SUPERNOVA EXPLOSIONS

- ~3 per century for a Milky Way type galaxy (Li et al. 2012).
- Liberate ~ 10^{58} neutrinos.
- Kinetic energies on the order of 10^{51} erg!
- Produced by stars with masses about 8 times more than the Sun, **massive stars**.

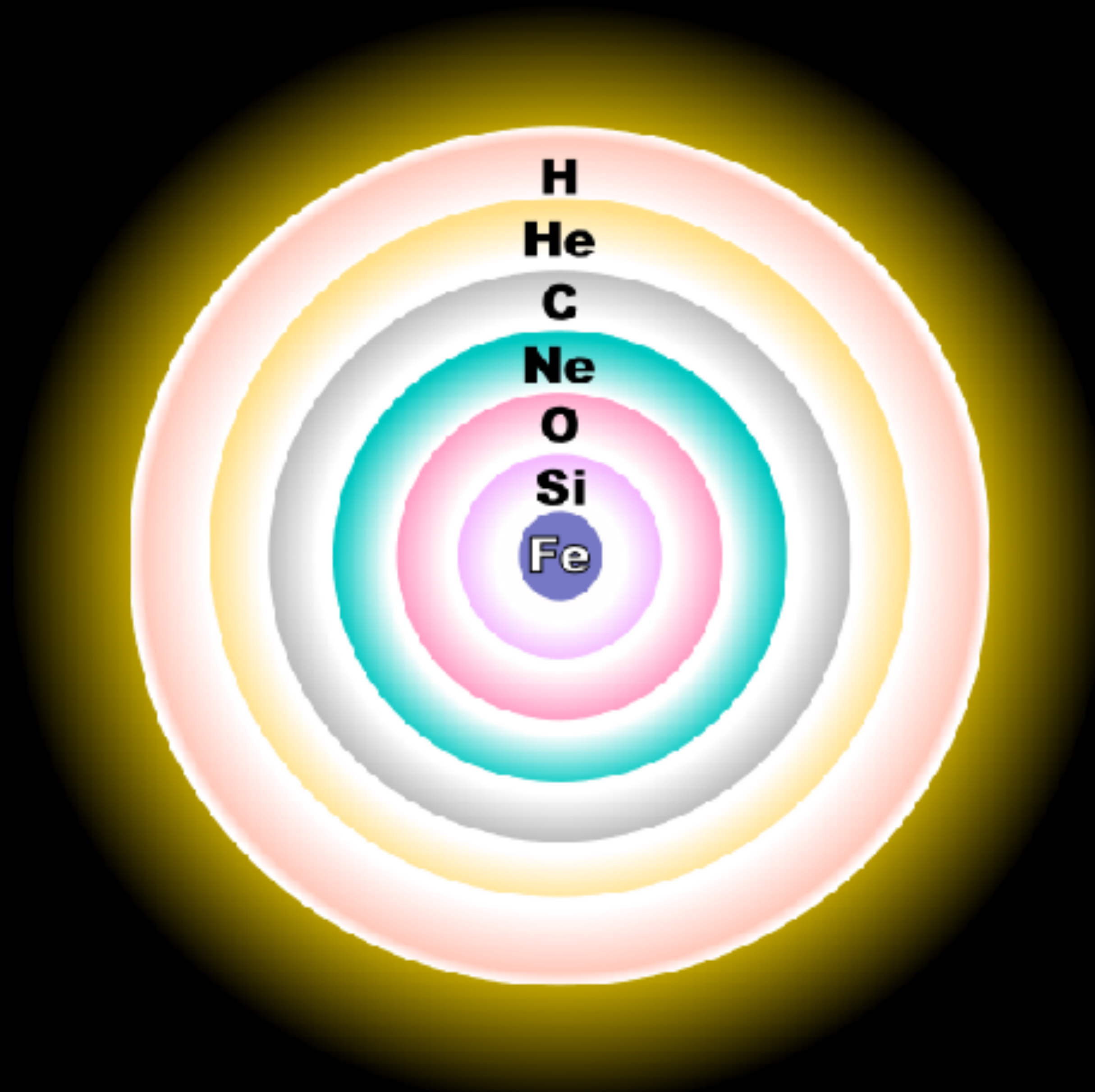


THE REMNANT OF SN 1987A. SOURCE: NASA GSFC.

CCSN Explosion Mechanism

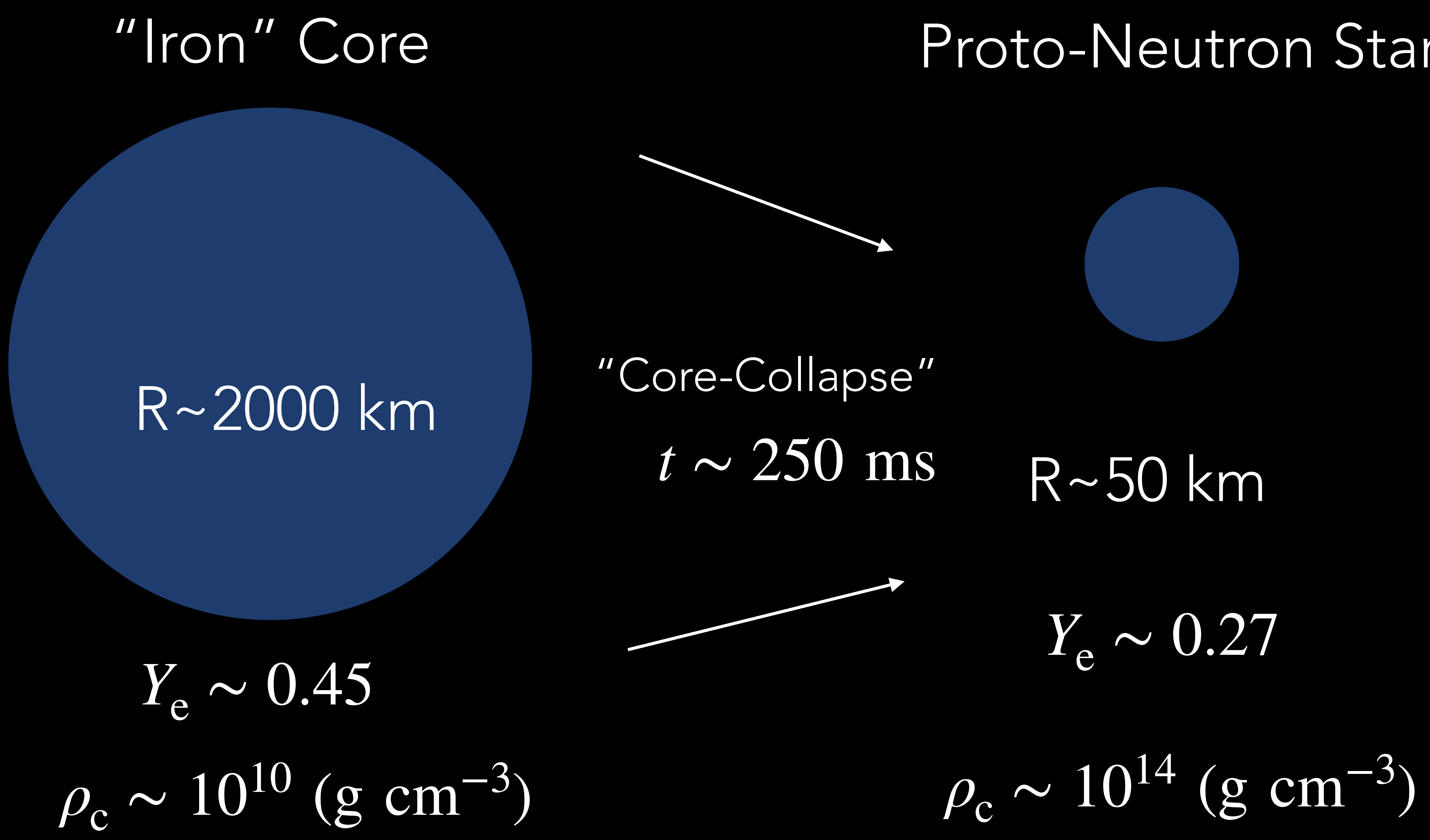
EVOLUTION TOWARDS IRON CORE-COLLAPSE IN A MASSIVE STAR

- Massive stars burn heavier and heavier elements.
- Form an inert core primarily of Fe peak elements.
- Core becomes gravitationally unstable as reactions remove pressure sources.
- Core collapses - rapidly !



CREDIT: R. J. HALL

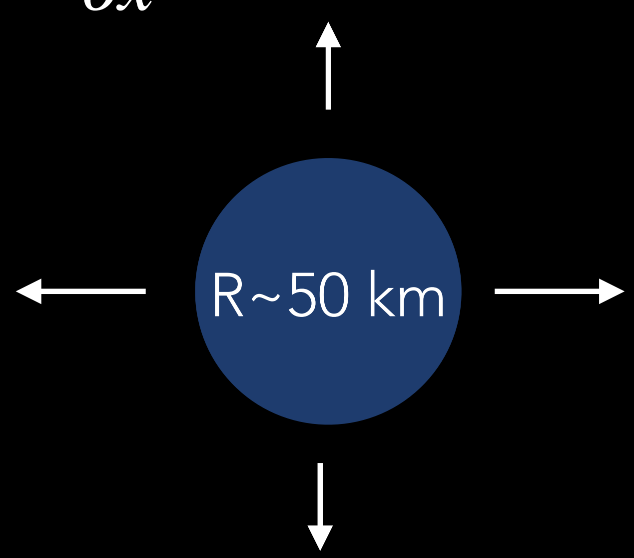
PHYSICS OF STELLAR CORE-COLLAPSE



PHYSICS OF STELLAR CORE-COLLAPSE

"Bounce"
Stiffening of Core
Launch Shock

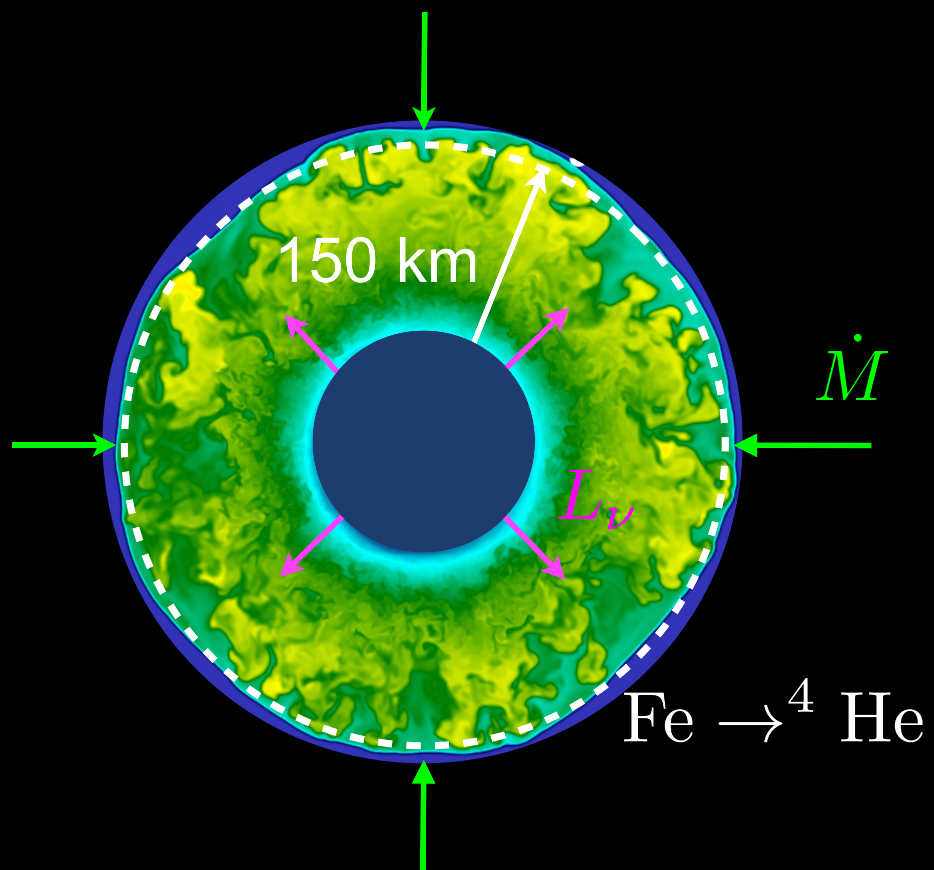
$$\frac{\partial}{\partial x}(P, v, \rho) \neq 0$$



"Bounce" to **Stalled** Shock

$t \sim 100$ ms

Stalled Shock



Entropy slice of explosion of 20 solar mass stars.
Credit: O' Connor & Couch (2018b).

Not enough energy to promptly explode star.

REVIVAL OF THE STALLED SHOCK

Delayed Neutrino Heating Mechanism

- Needs ~10^{51} erg to unbind the star, explode.
- PNS contraction releases energy as neutrinos ~ 10^{53} erg / s !!
- Heating by neutrinos beneath the stalled shock via absorption.
- *Only* need a few % of released neutrinos to drive explosion (Bethe & Wilson 1985).

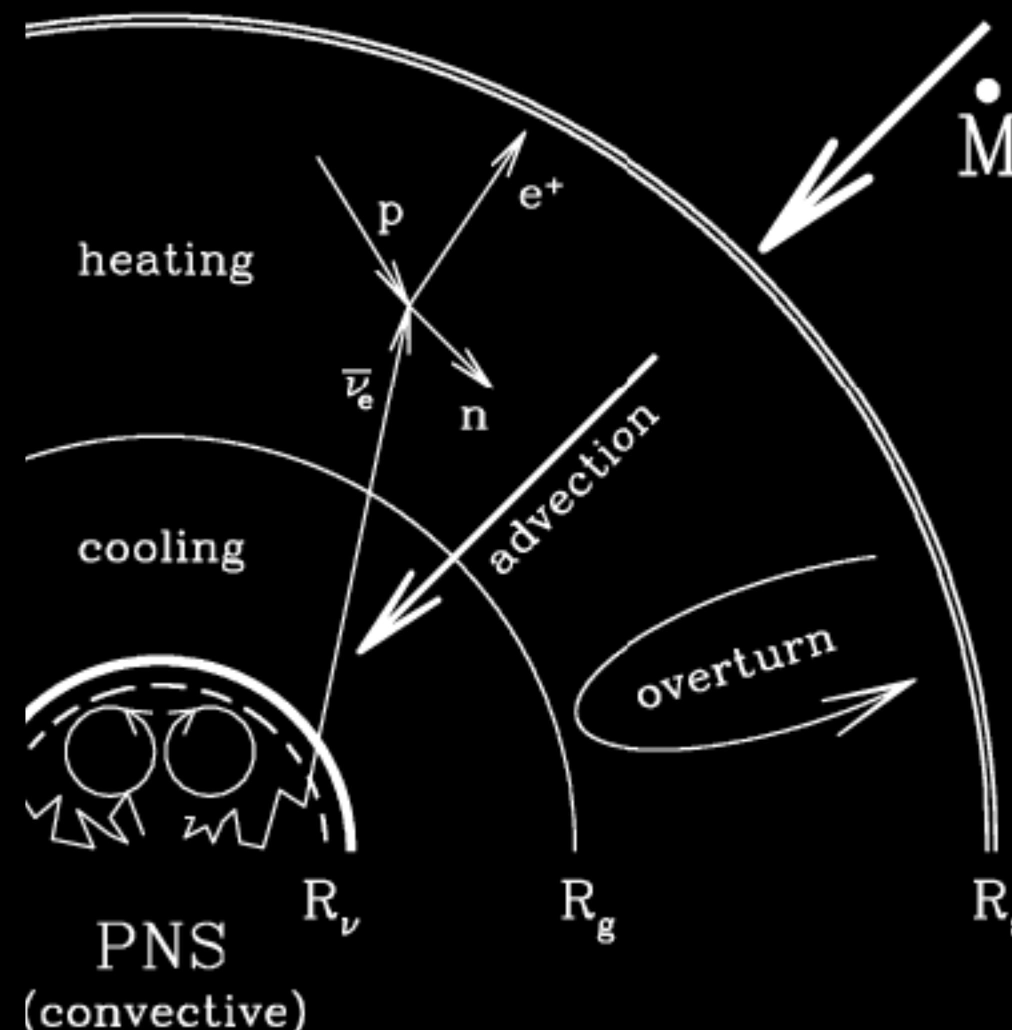


Diagram showing revival of stalled shock.
Credit: Janka (2011).

ERA OF 3D CCSN SIMULATIONS

<u>Fully-coupled!</u>

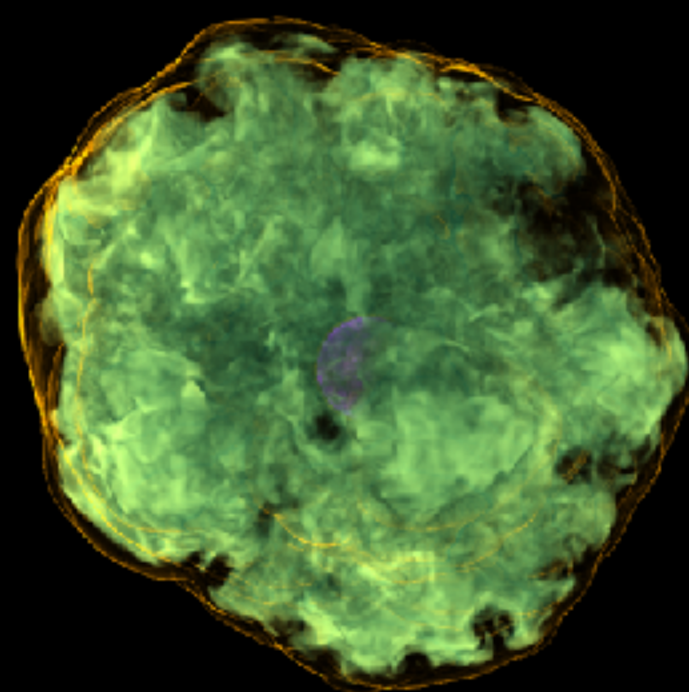
3D Magnetohydrodynamics

General Relativity

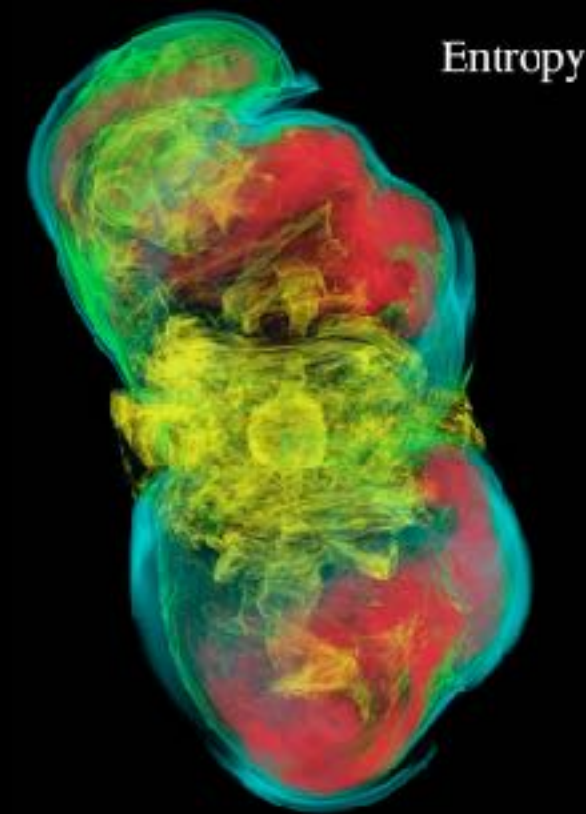
Boltzmann ν-transport

Microphysics
(Nuclear EOS, ν-interactions, nuclear kinetics)

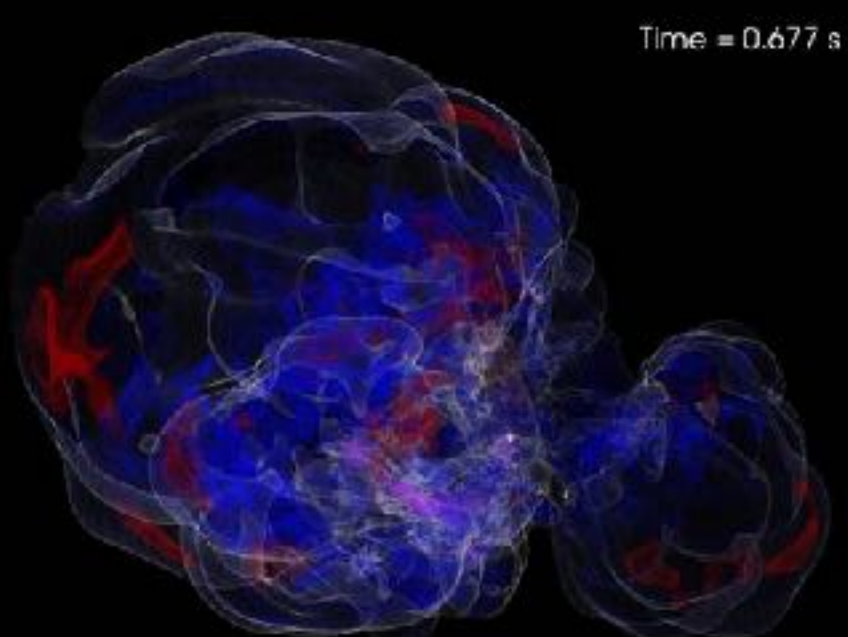
Credit: Sean Couch



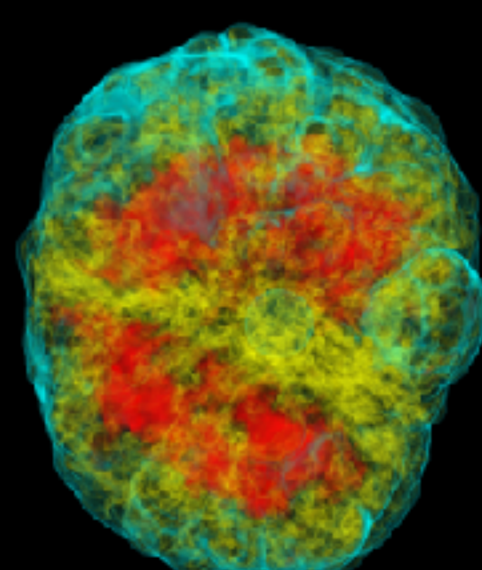
(Fields + 2022b, in prep.)



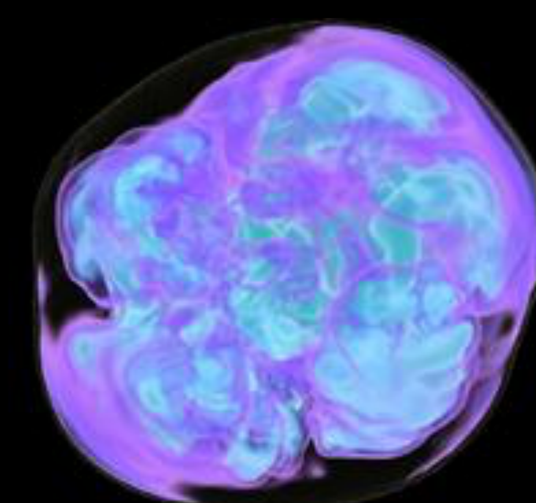
(Moesta + 2014)



(Vartanyan+ 2019)



(Roberts + 2016)



(Burrows + 2019)

Solved problem…right?

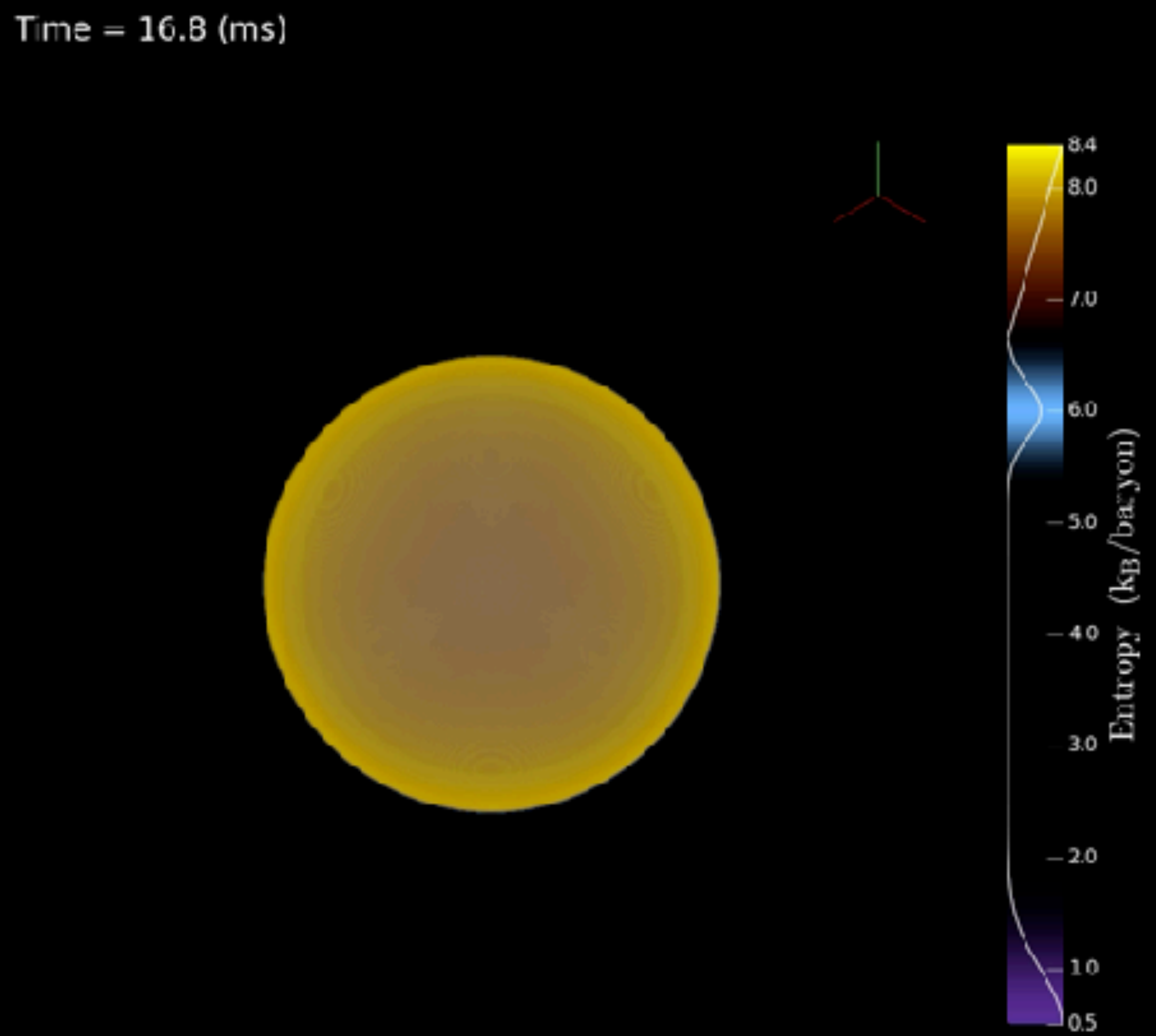
INTRODUCTION

The CCSN “Problem” and possible solutions

THE CORE-COLLAPSE 'PROBLEM'

How do we (try) to model stellar explosions?

- 1D Stellar Evolution Codes for pre-supernova evolution.
- Evolve explosion in 2/3D using multi-D hydro codes.
- Shock failed to be revived in some models.

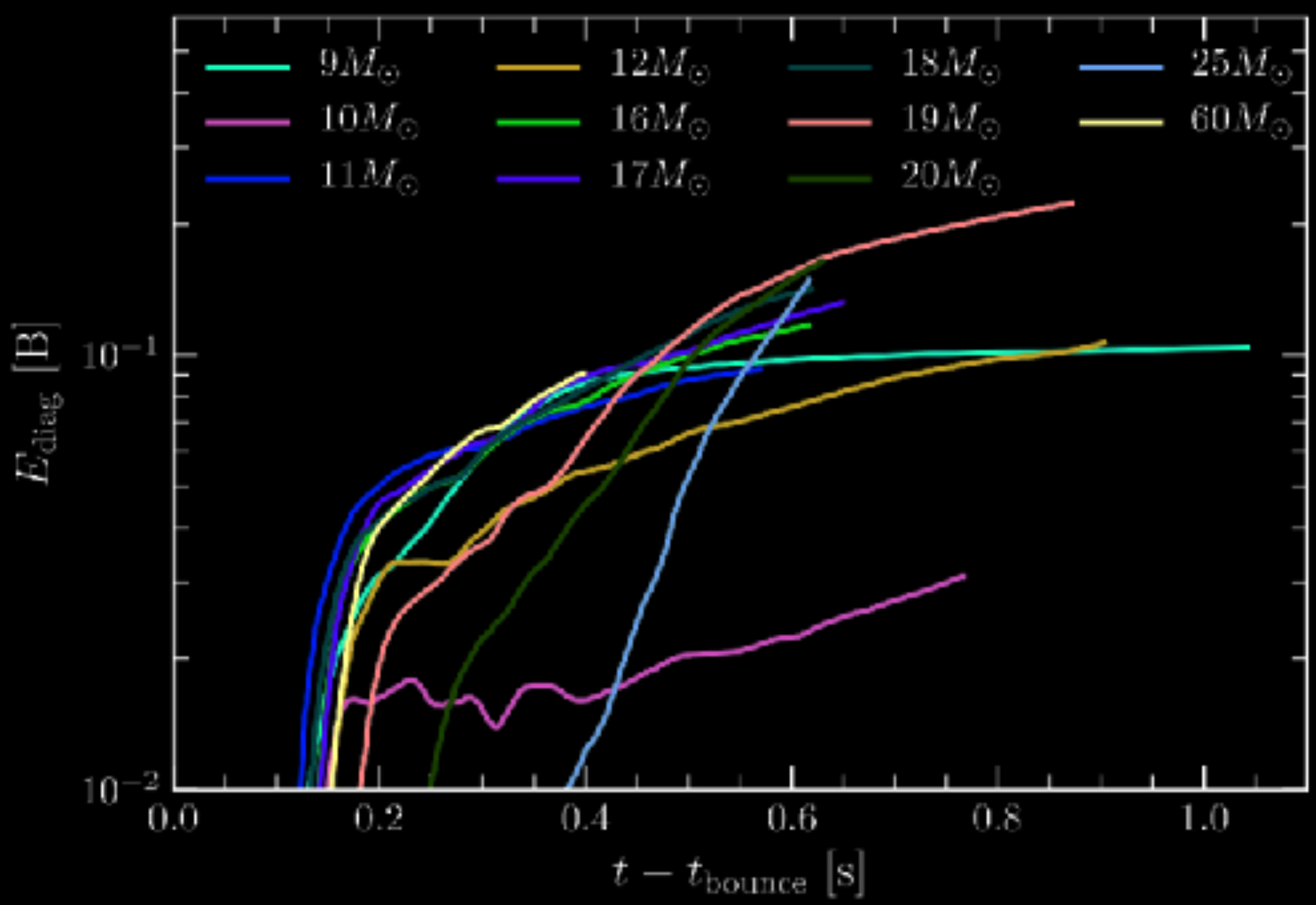


Failed explosion using spherically symmetric 1D model from Couch + 2018.

THE CORE-COLLAPSE 'PROBLEM'

How do we (try) to model stellar explosions?

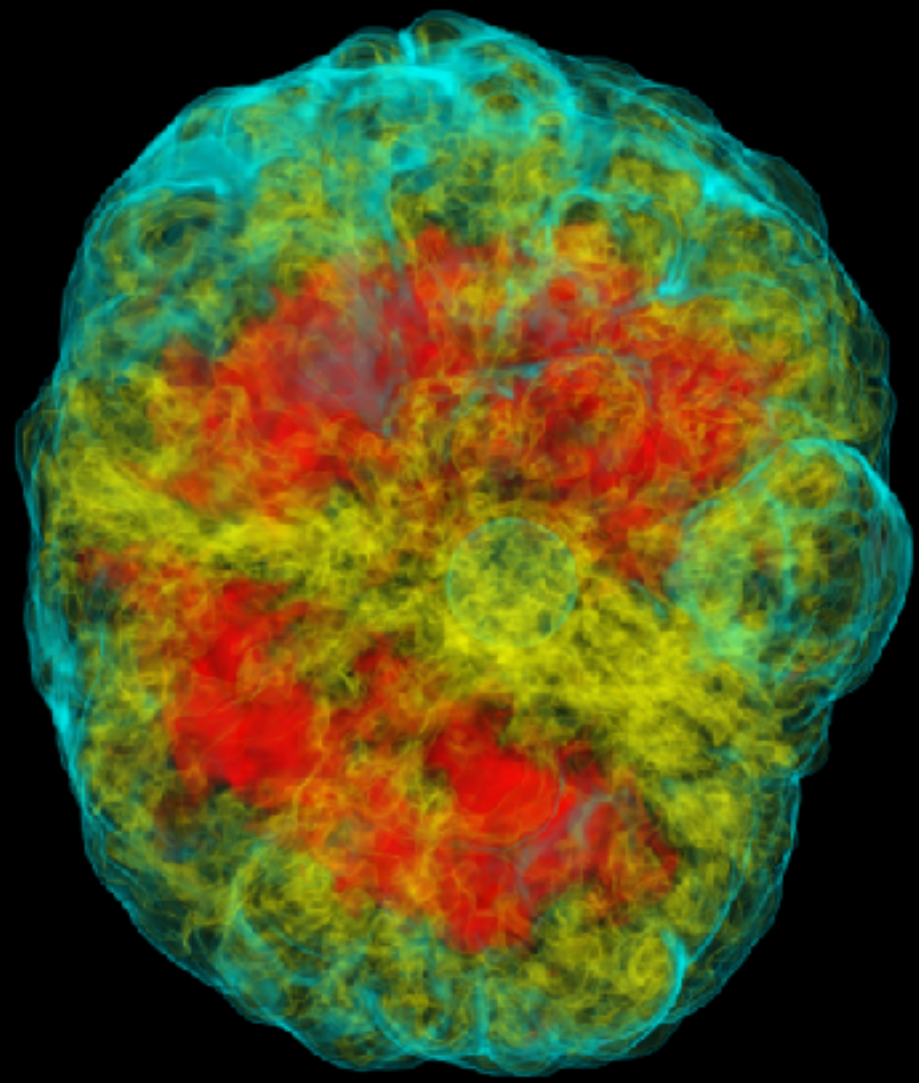
- Struggle to match range of Type IIP explosion energies of ~0.5-4B *(Kasen & Woosely 2015)*.
- 3D exploding models show low energies?
- Need to reach asymptotic plateau requires longer simulations *(Burrows+ 2019)*.



Evolution of explosion energy for 3D CCSN models from Burrows + 2019.

SOLUTION(S) TO THE CORE-COLLAPSE ‘PROBLEM’?

- General Relativistic Gravity - More compact PNs lead to larger neutrino luminosities.
- Sophisticated Neutrino Transport - Full Transport + GR can result in explosion.
- Initial models/Perturbations - Pre-SN models are **not** spherical and can vary.

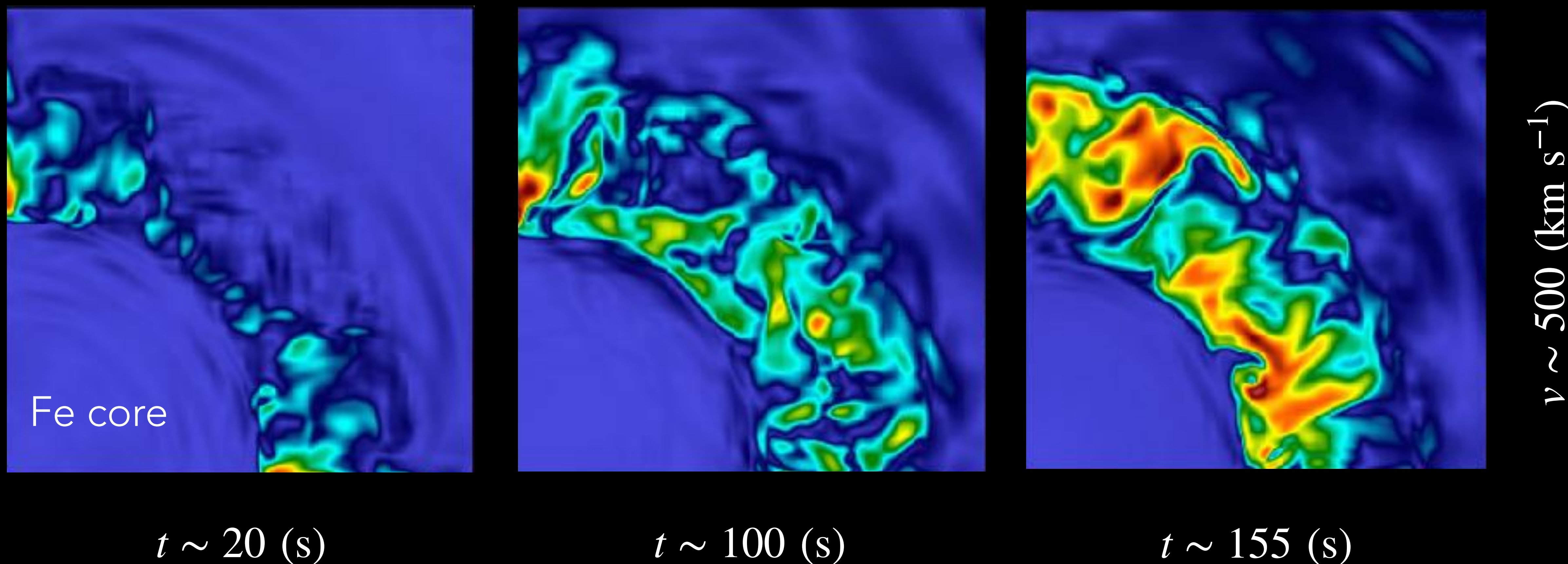


Volume rendering of the entropy distribution from *Roberts + 2016*.

INTRODUCTION

Deeper look in to the Pre-Supernova Models

PERTURBATIONS IN THE PRE-SUPERNOVA MODEL

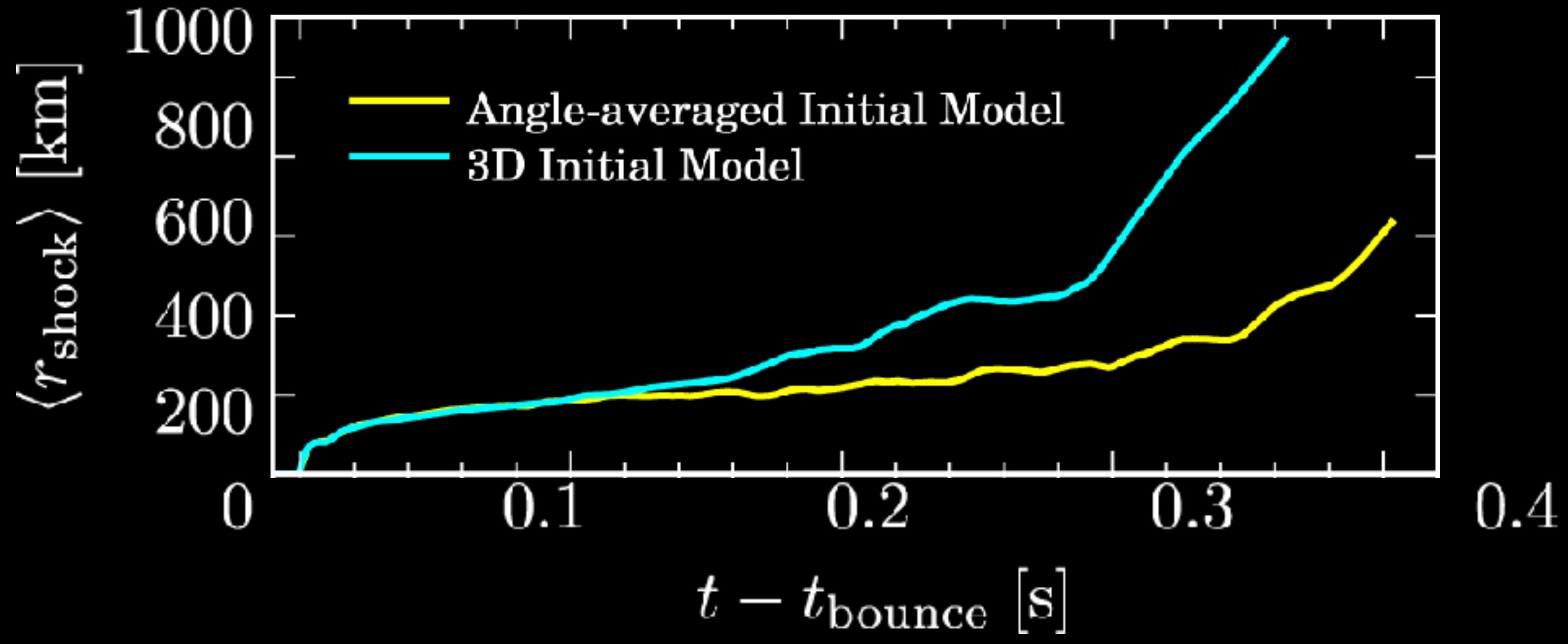


(Couch + ApJL, 2015)

- 3D Octant model, ~ three minutes, evolved using 21 isotope network.

PERTURBATIONS IN THE PRE-SUPERNOVA MODEL

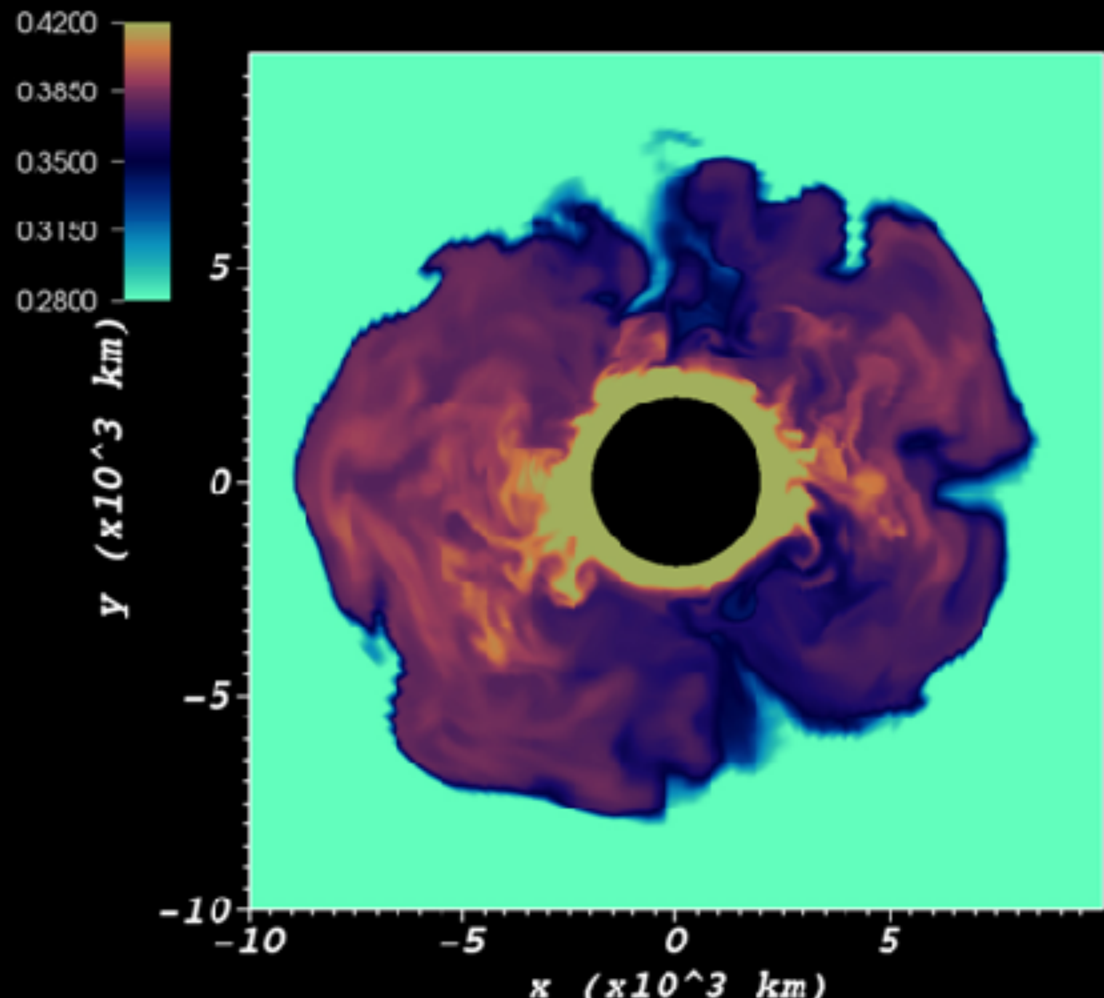
3D Initial model leads to faster, stronger explosion.



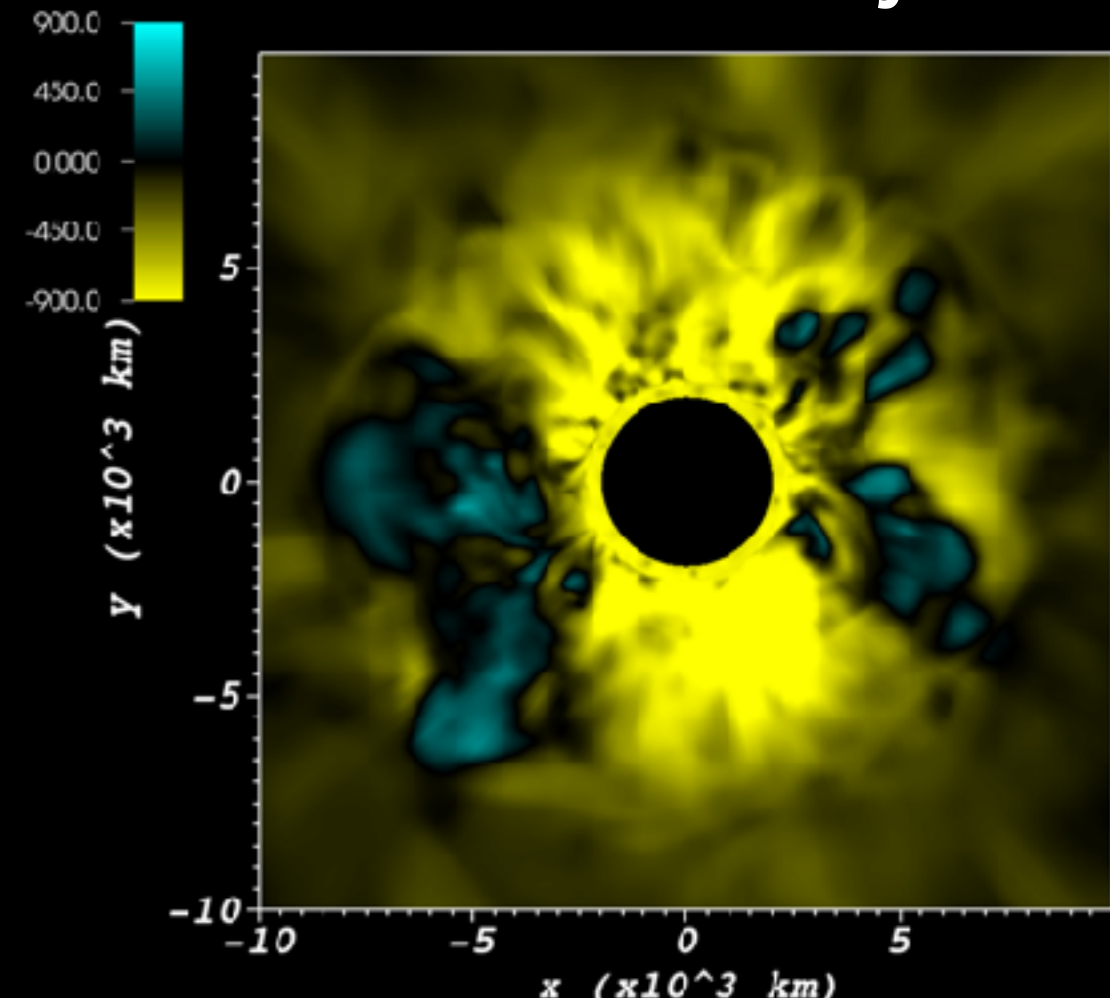
- Multi-D progenitors provide a solution to the core-collapse problem.

MULTI-DIMENSIONAL SIMULATIONS OF MASSIVE STARS

Silicon-28

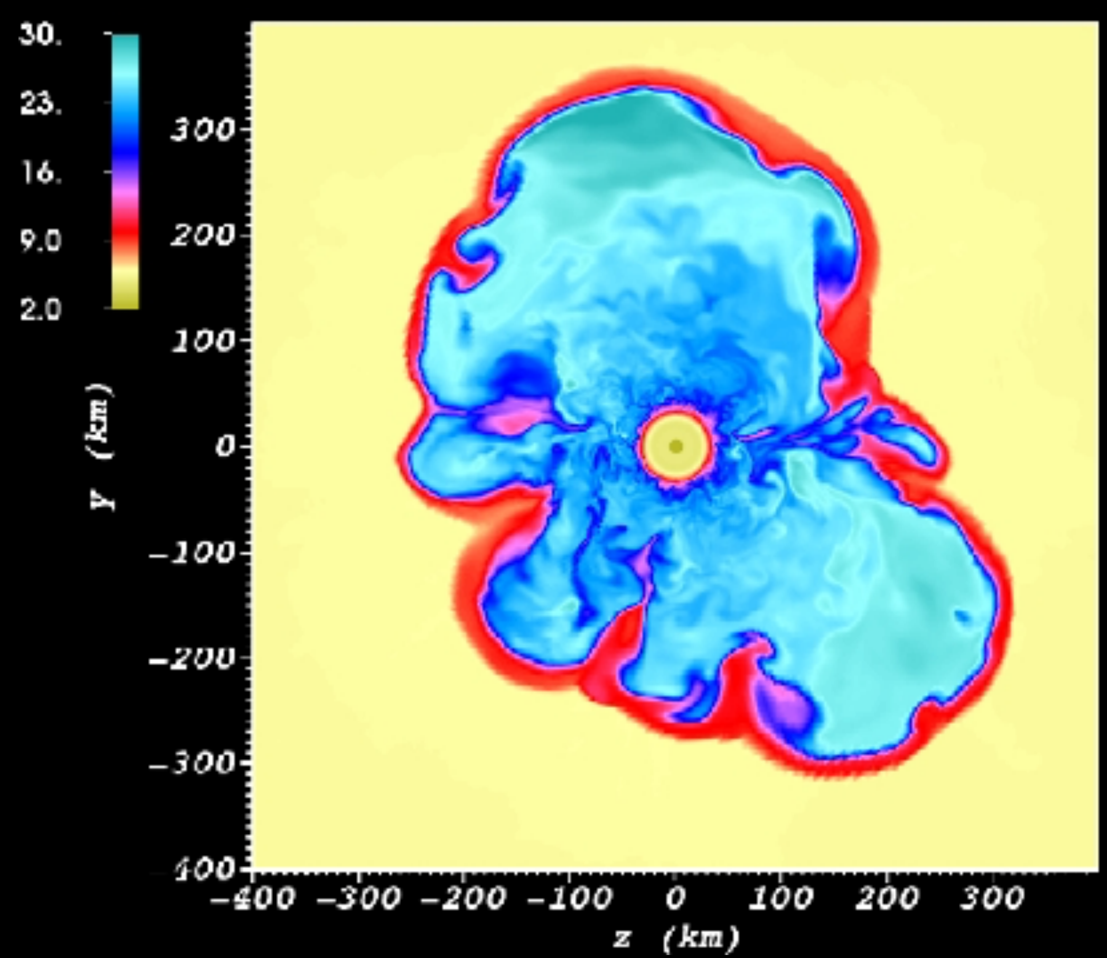


Radial Velocity

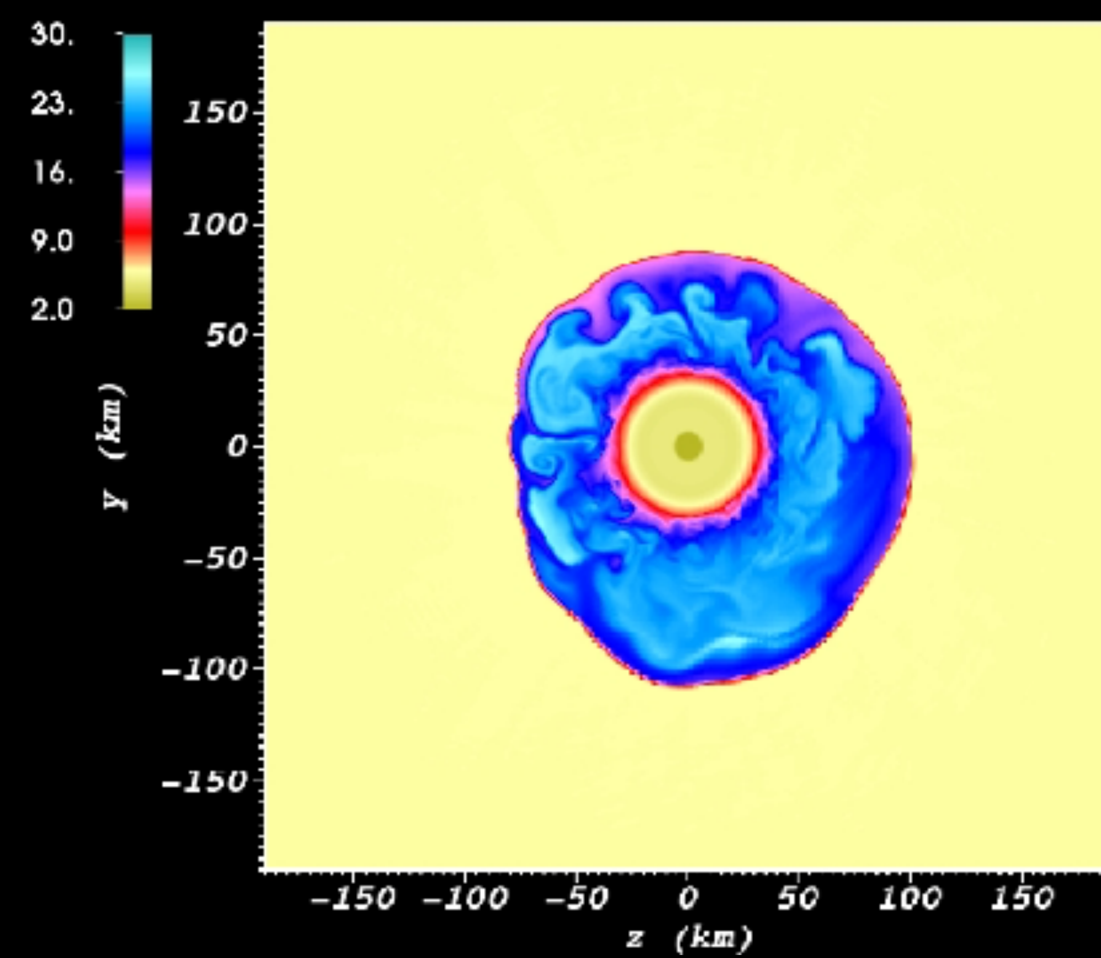


- 4pi simulations of oxygen shell burning find bipolar flow near collapse in simulation of 18 solar mass star. *(Muller +2016)*

IMPACT OF PROGENITORS ON EXPLOSION MECHANISM



3D initial progenitor



1D initial progenitor

(Muller + 2017)

IMPACT OF PROGENITORS ON EXPLOSION MECHANISM

How do 3D progenitors help facilitate explosion?

- **Large mach numbers** cause density fluctuations favorable for explosion.

$$\delta\rho/\rho \propto \mathscr{M}_{\mathrm{prog.}}$$

- **Increase mass in gain** region due to non-radial flow in post-shock region.

$$\dot{Q}_\nu \propto M_{\mathrm{gain}}$$

(Muller + 2017)

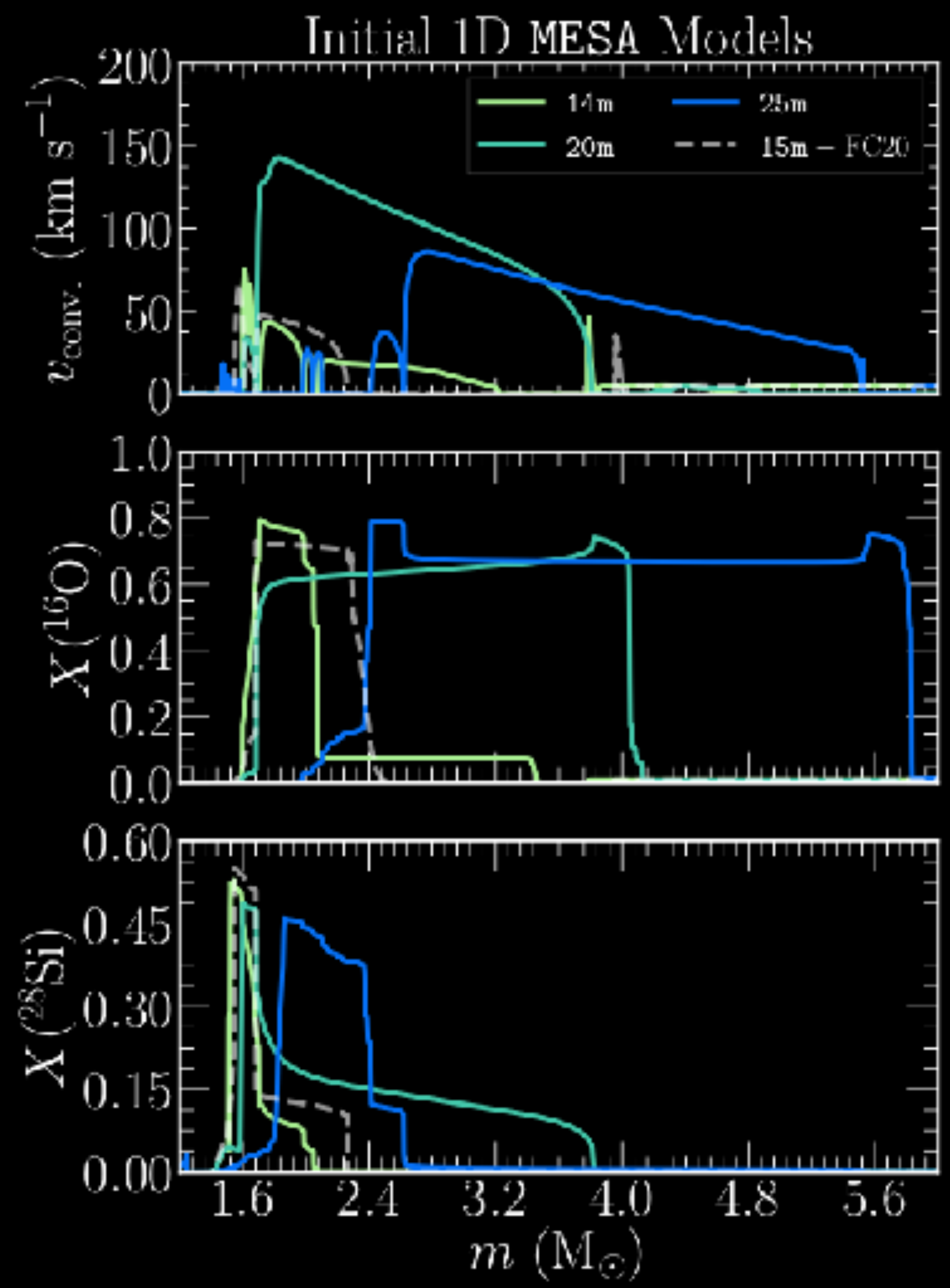
- Increase in non-radial kinetic energy at **large** scales.

(Couch + 2014, 2015)

Convection in multiple 3D Progenitor Models

MASSIVE STAR CONVECTION IN MULTIPLE PROGENITORS

- 3D simulations using FLASH for 14-,20-, and 25 $M_{\odot}$ models.
- Evolved ~**10 minutes** collapse using approximate network.

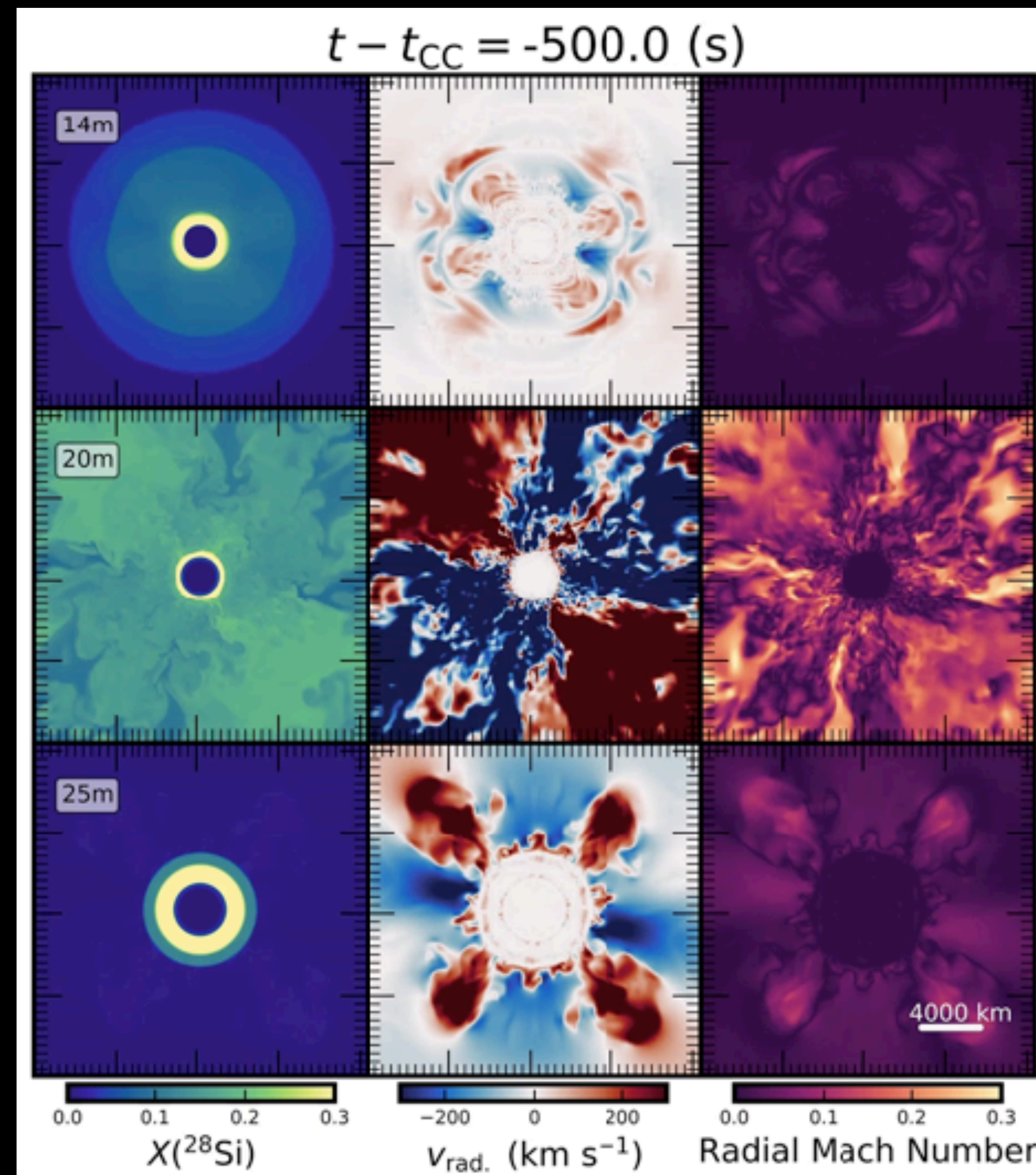


Initial 1D profile structure for 3D models.
(Fields & Couch 2021)

MASSIVE STAR CONVECTION IN MULTIPLE PROGENITORS

- Models vary in convective speeds!
- Large-scale flow observed in 20 $M_\odot$ model.

$$\delta\rho/\rho \propto \mathscr{M}_{\rm prog.}$$



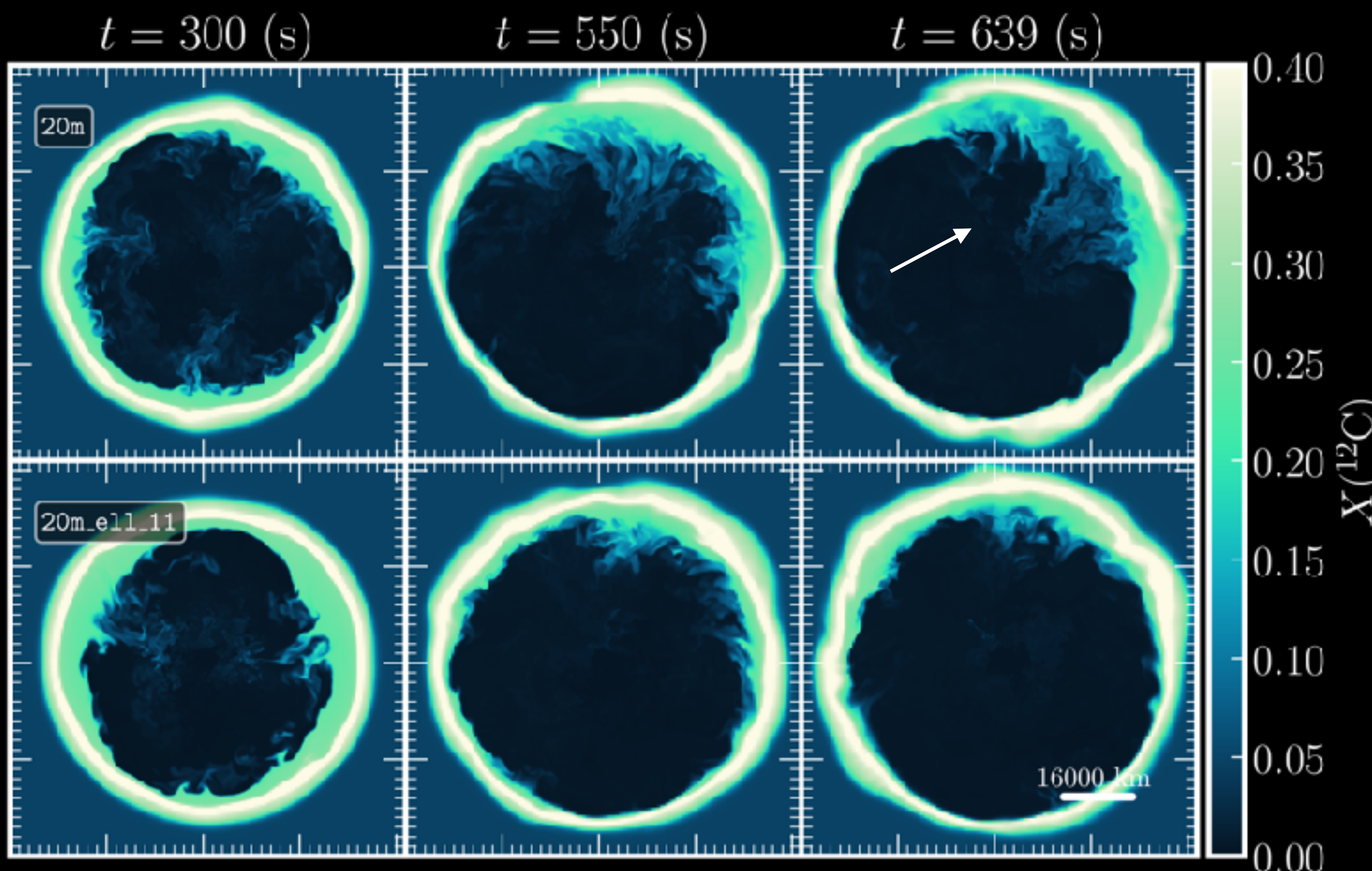
SIMULATIONS OF MASSIVE STAR CONVECTION IN MULTIPLE PROGENITORS

- Smaller O-shell Region, smaller mach numbers, ~0.04!
- Convection occurring at broad range of scales.

$$M_{\rm ZAMS} = 14 M_\odot$$
$$t - t_{cc} = -300\ (s)$$

Volume rendering of the velocity field for 3D progenitor model near collapse ***(Fields & Couch 2021a.)***.

MASSIVE STAR CONVECTION IN MULTIPLE PROGENITORS



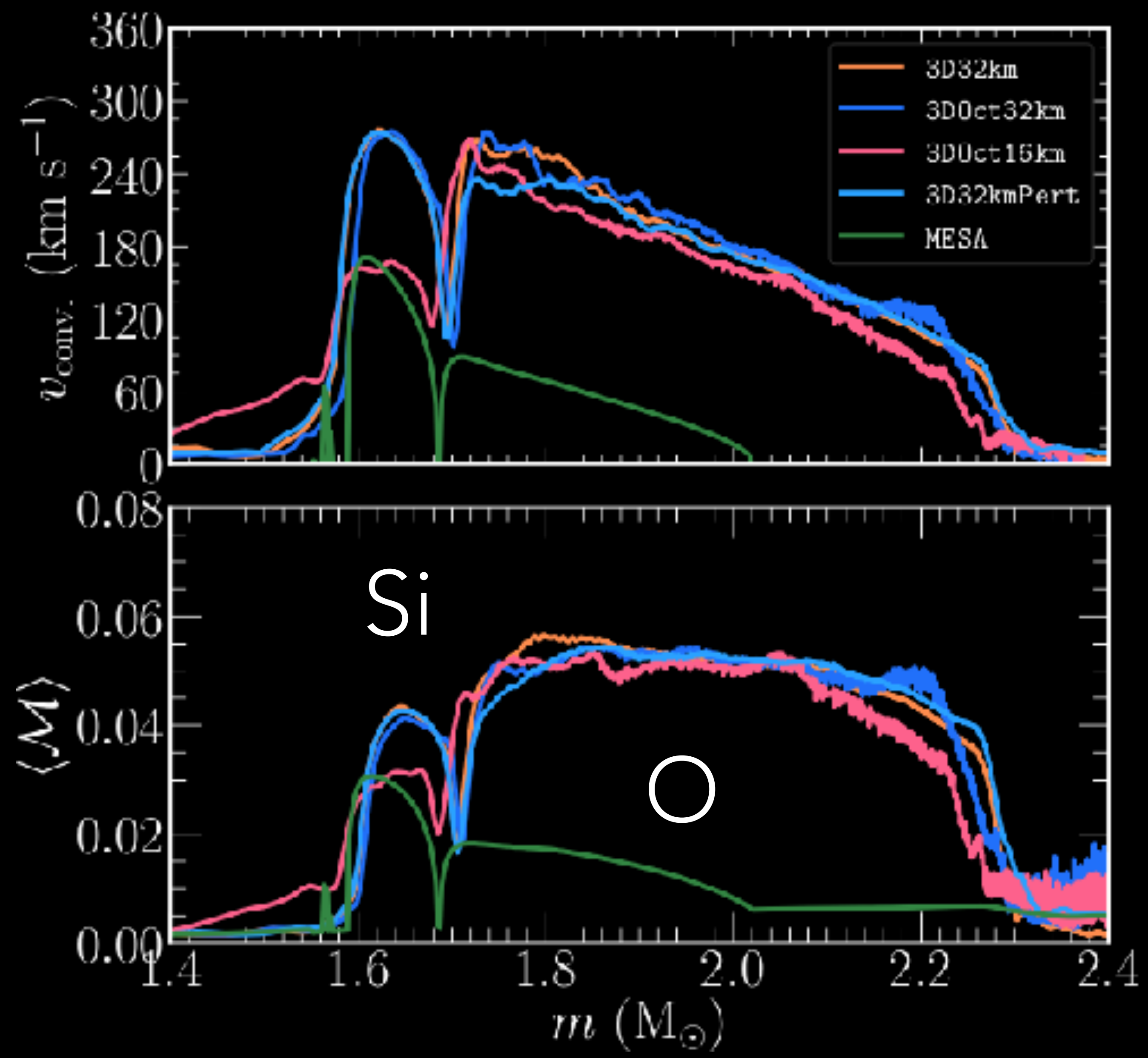
$(5.2 - 7.5) \times 10^{-4} M_\odot \ \mathrm{yr}^{-1}$

(Fields & Couch 2021a.).

C-ingestion in the O-shell region affected by initial perturbations.

MULTI-DIMENSIONAL SIMULATIONS OF MASSIVE STARS

- 1D MESA model matches Si-shell convection well.
- Largely under predicts O-shell speeds and extent.
- 1D approximation good, in some cases.



Angle average mach number profiles for all models at different times ***(Fields & Couch 2020)***.

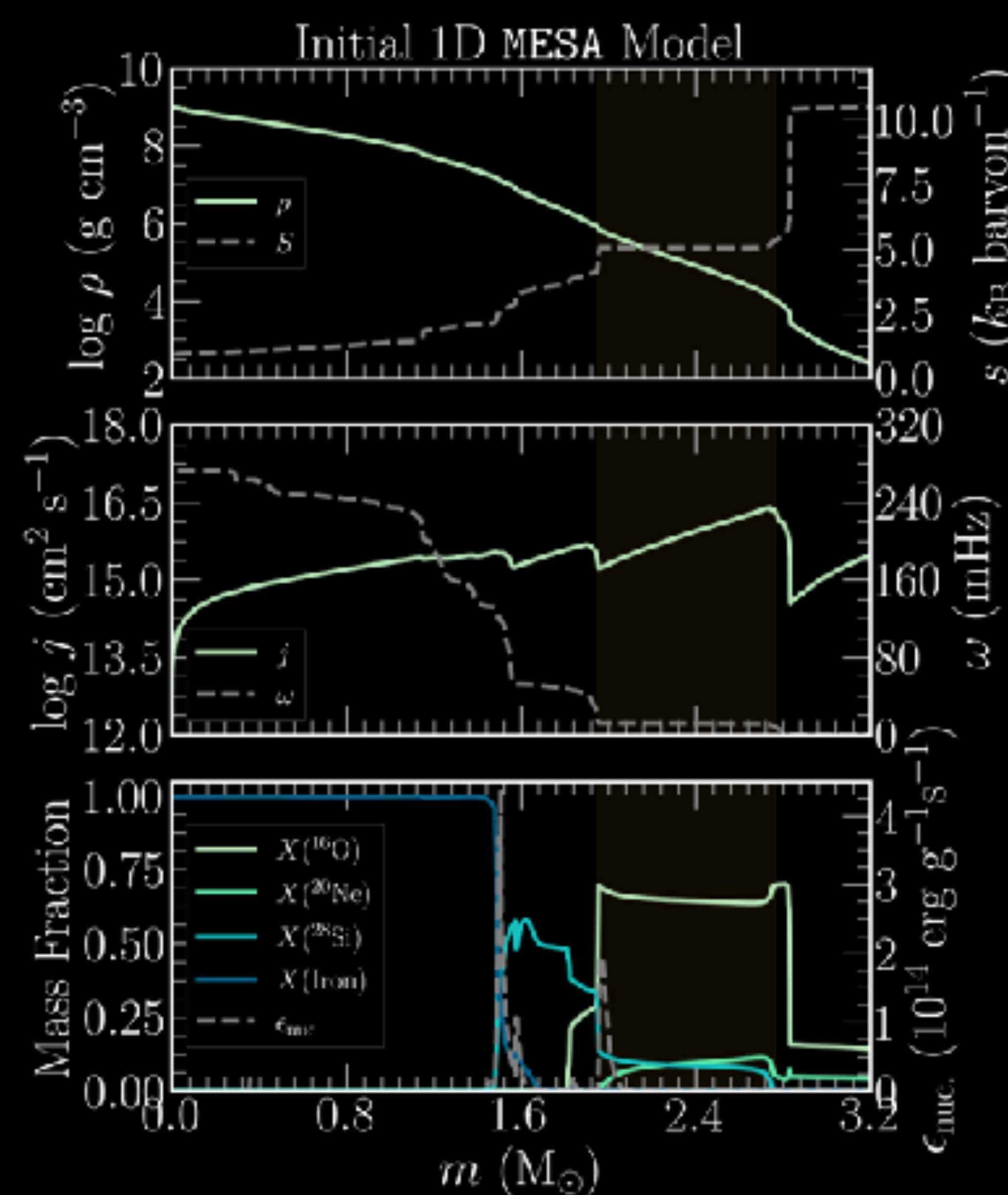
3D CCSN PROGENITORS

3D Evolution of a Rapidly Rotating $16 M_\odot$ Star

(… what about rotation?)

CONVECTION IN RAPIDLY ROTATING PROGENITORS

- 3D simulations using FLASH for $16M_{\odot}$ model.
- Rotation initialized to 350 km/s at ZAMS.
- Evolved the final 10 minutes to iron core-collapse.
- Includes complete iron core.

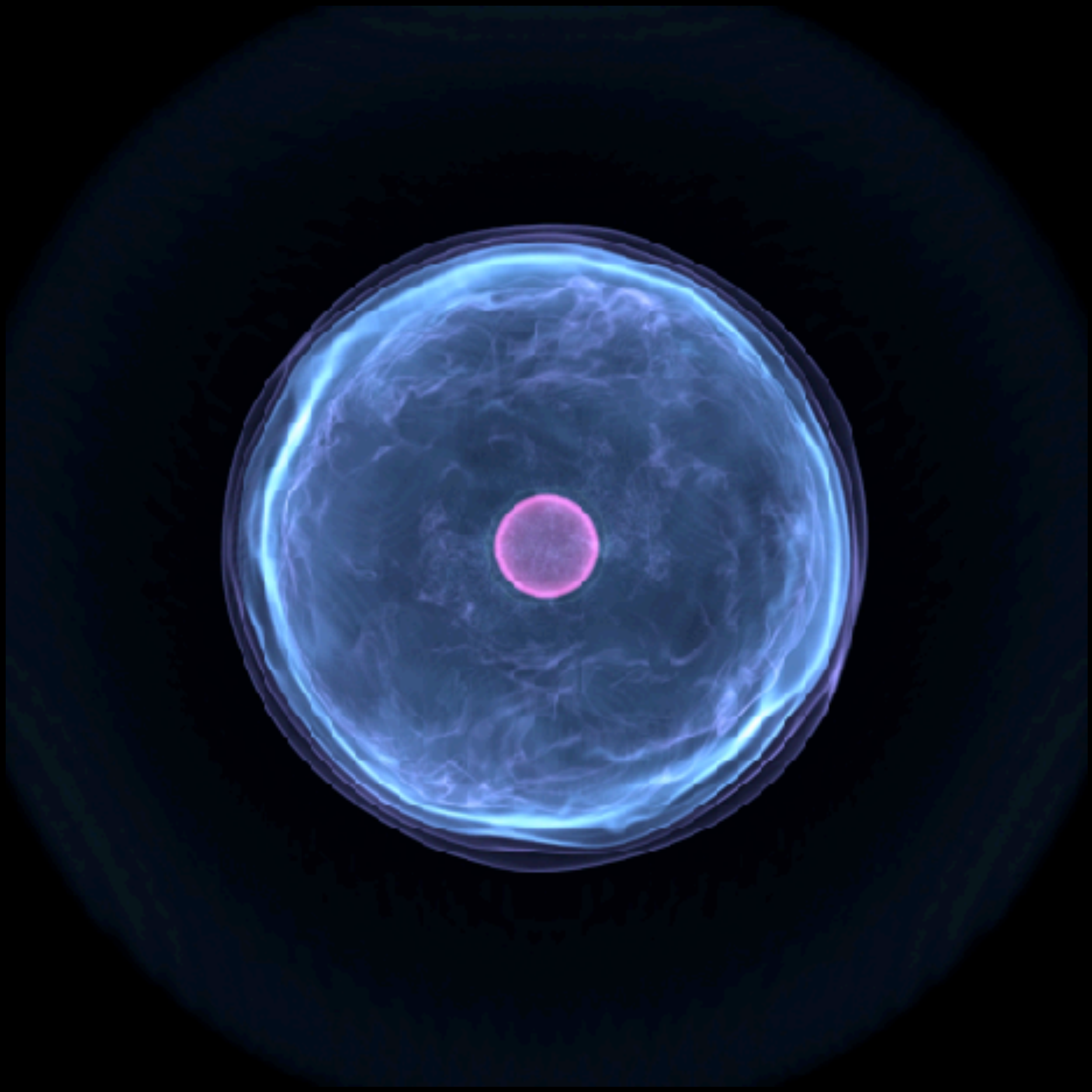


Initial 1D profile structure for 3D model.

(Fields, 2022)

MASSIVE STAR CONVECTION IN ROTATING PROGENITORS

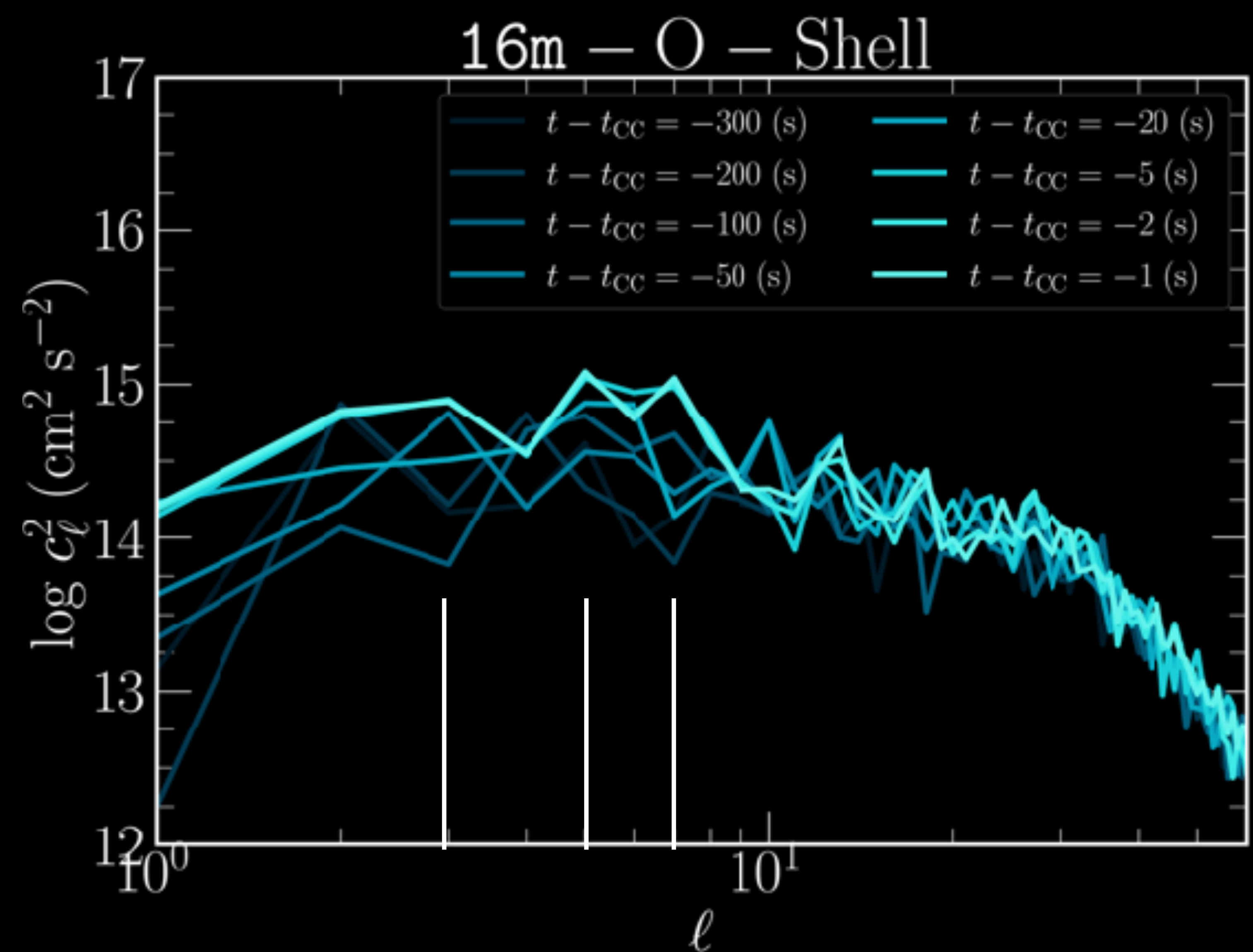
- Broad convective scales
- Relatively weak Mach numbers ~0.04.
- Weak Si-shell convection.



Volume rendering of the Ne-20 mass fraction.
(Fields 2022)

MASSIVE STAR CONVECTION IN ROTATING PROGENITORS

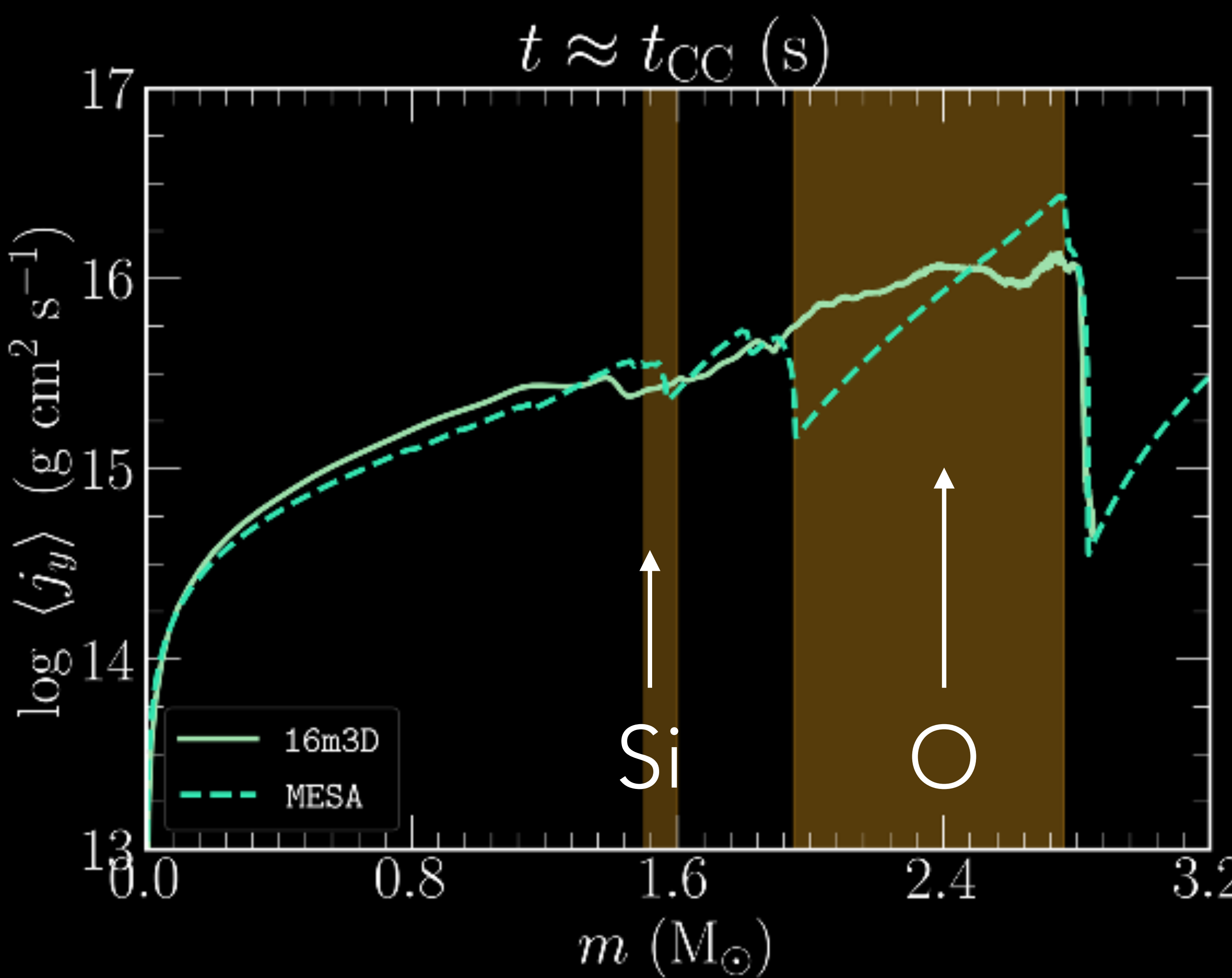
- Convection across a range of scales.
- Flow tends towards large scales at late times ($\ell = 3,5,7$).



Spectrum of radial velocity field for 3D rotating progenitor.
(Fields 2022)

MASSIVE STAR CONVECTION IN ROTATING PROGENITORS

- AM profile diverges from MESA in convective regions.
- We find a NS spin period of $P \sim 1.42$ (ms) at collapse.
- MESA model finds $P \sim 1.41$ (ms).

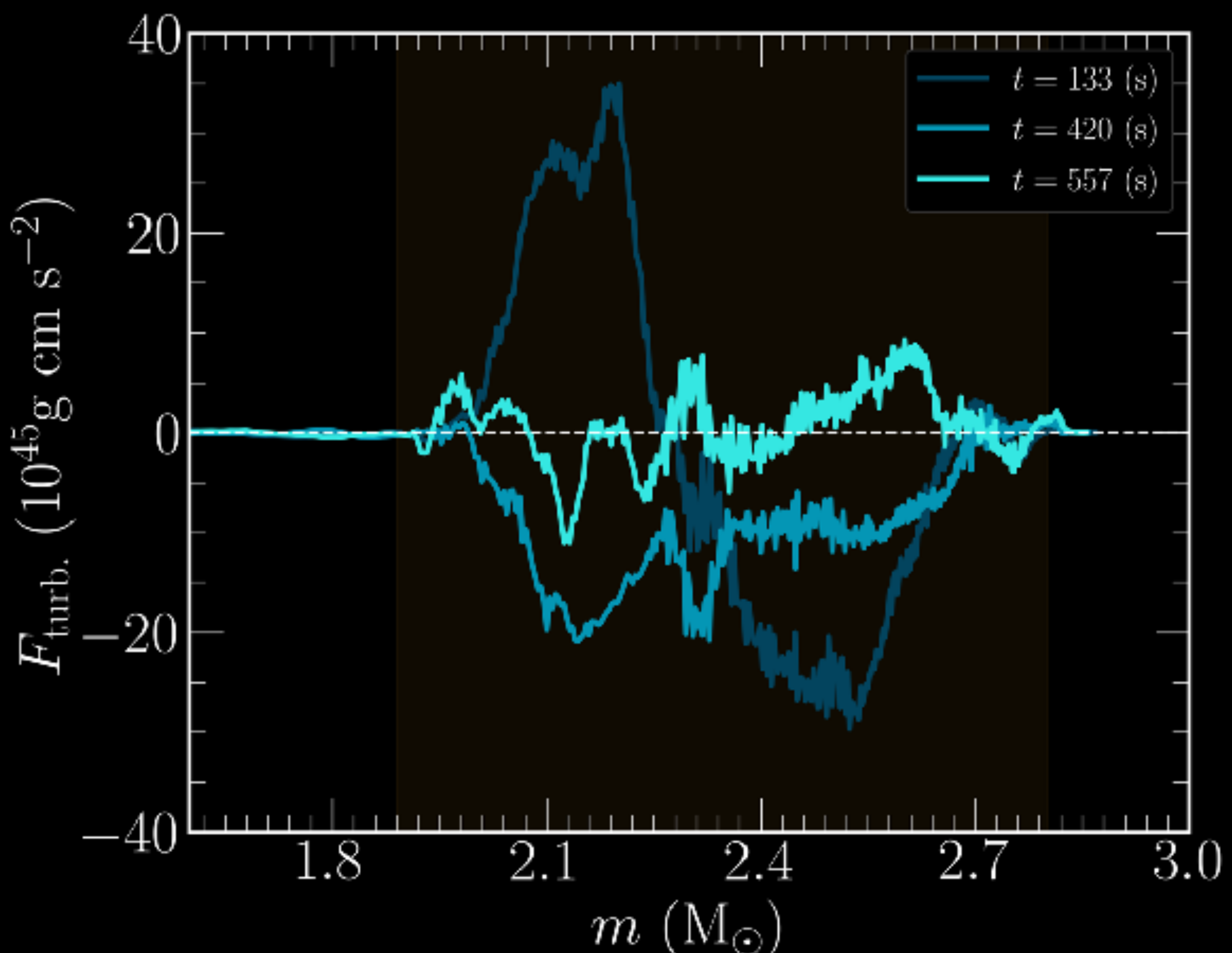


Angular momentum profiles for rotating 3D progenitor.
(Fields 2022)

MASSIVE STAR CONVECTION IN ROTATING PROGENITORS

- Advective term in non-convective regions.
- Angular momentum flux components.
- Positive flux in the O-shell.

$$F_{\text{turb.}} = \left\langle \rho v_r'' j_y'' \right\rangle$$



Angular momentum flux profiles.

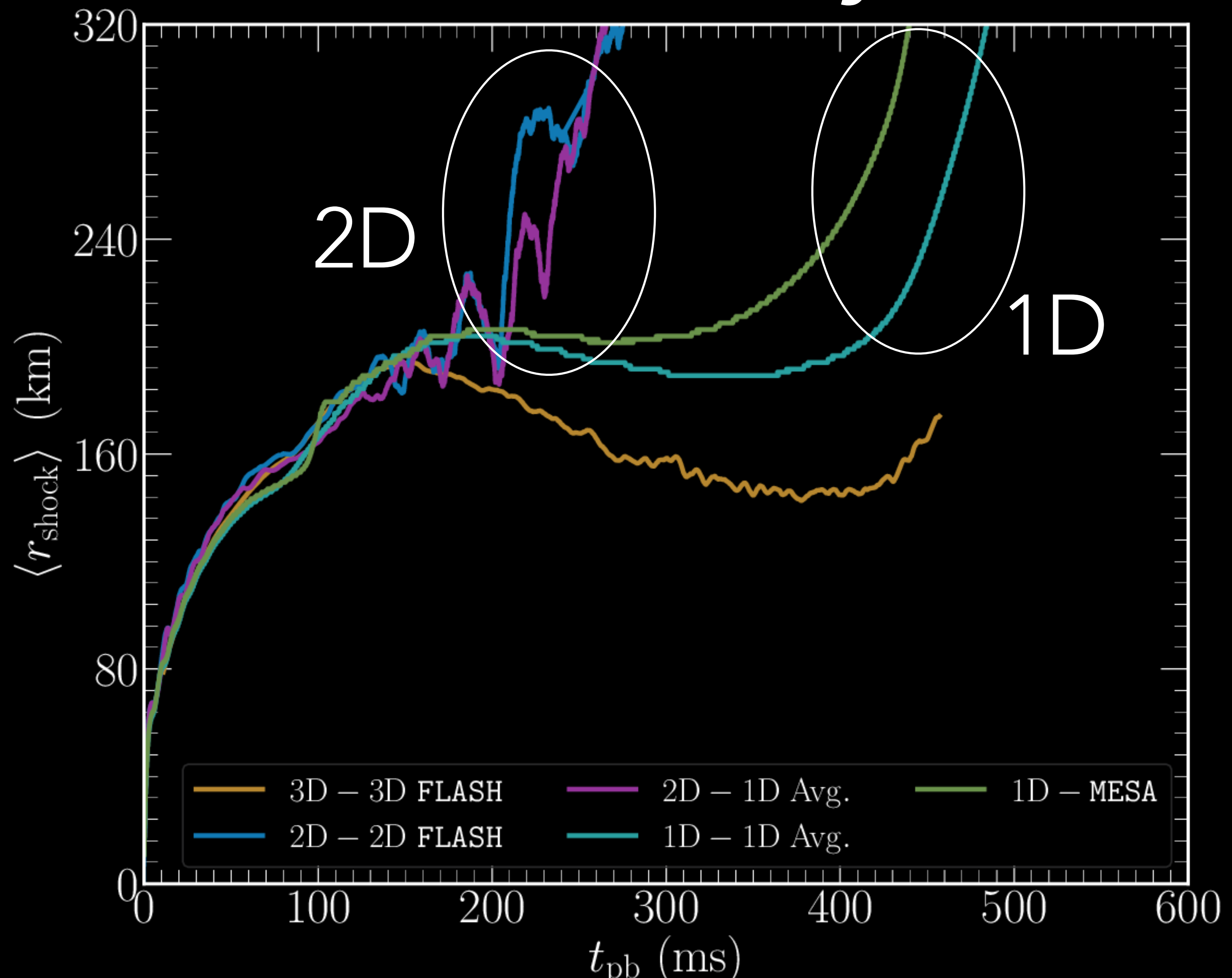
(Fields 2022)

CCSNe using 3D Progenitors

CCSN EXPLOSIONS OF MULTI-D PROGENITORS

- 1/2/3D CCSN simulations.
- Use 2D/3D progenitors.
- Multi-group/species, energy/velocity dependent neutrino transport, **M1**.

Preliminary



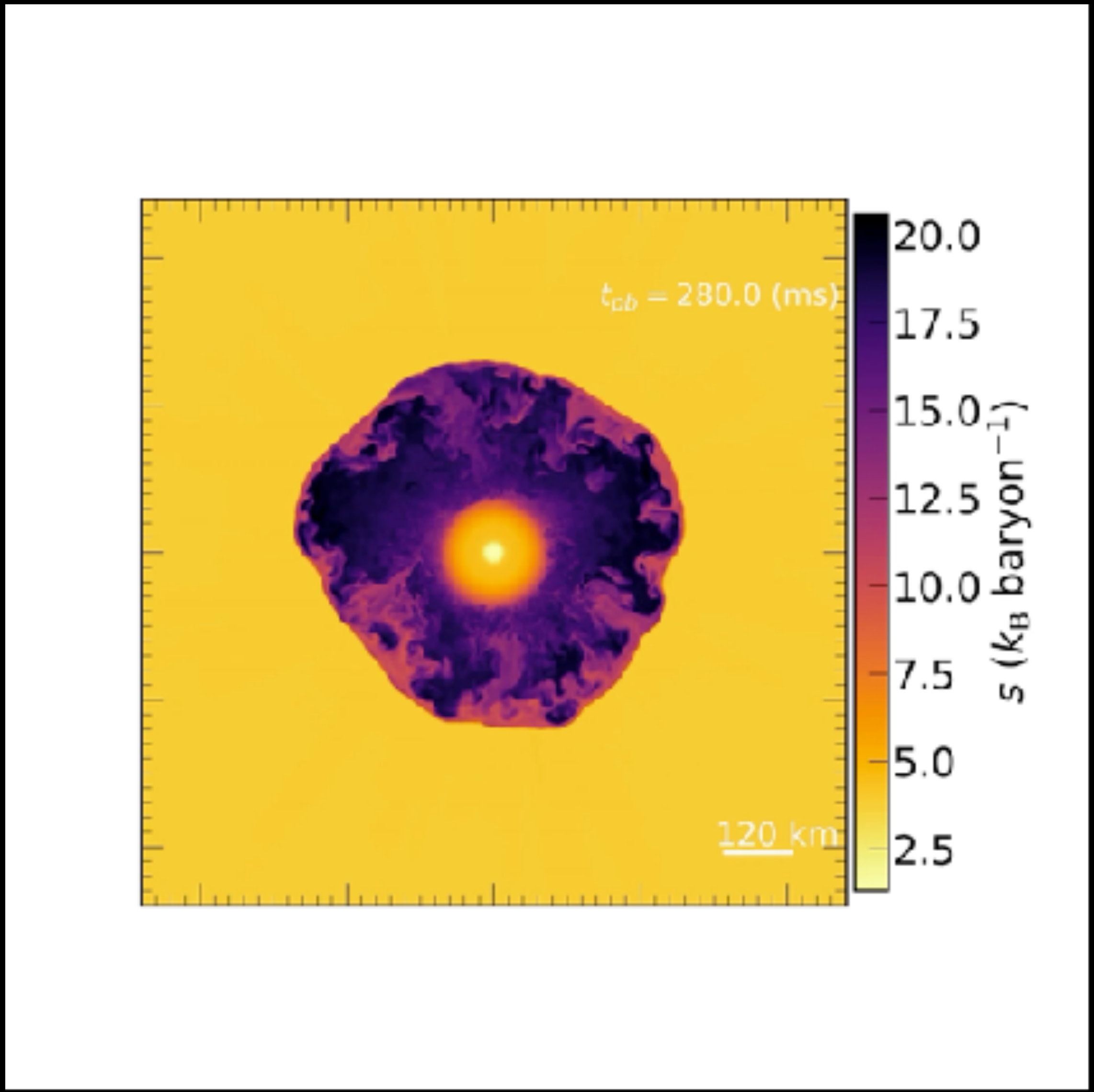
Mean shock radius evolution for multi-D CCSN models

(Fields + 2022b, in prep.).

CCSN EXPLOSIONS OF MULTI-D PROGENITORS

- 3D model approaching shock runaway.
- Large non-radial kinetic energy.
- Test for LESA, implications for NS kick, etc.

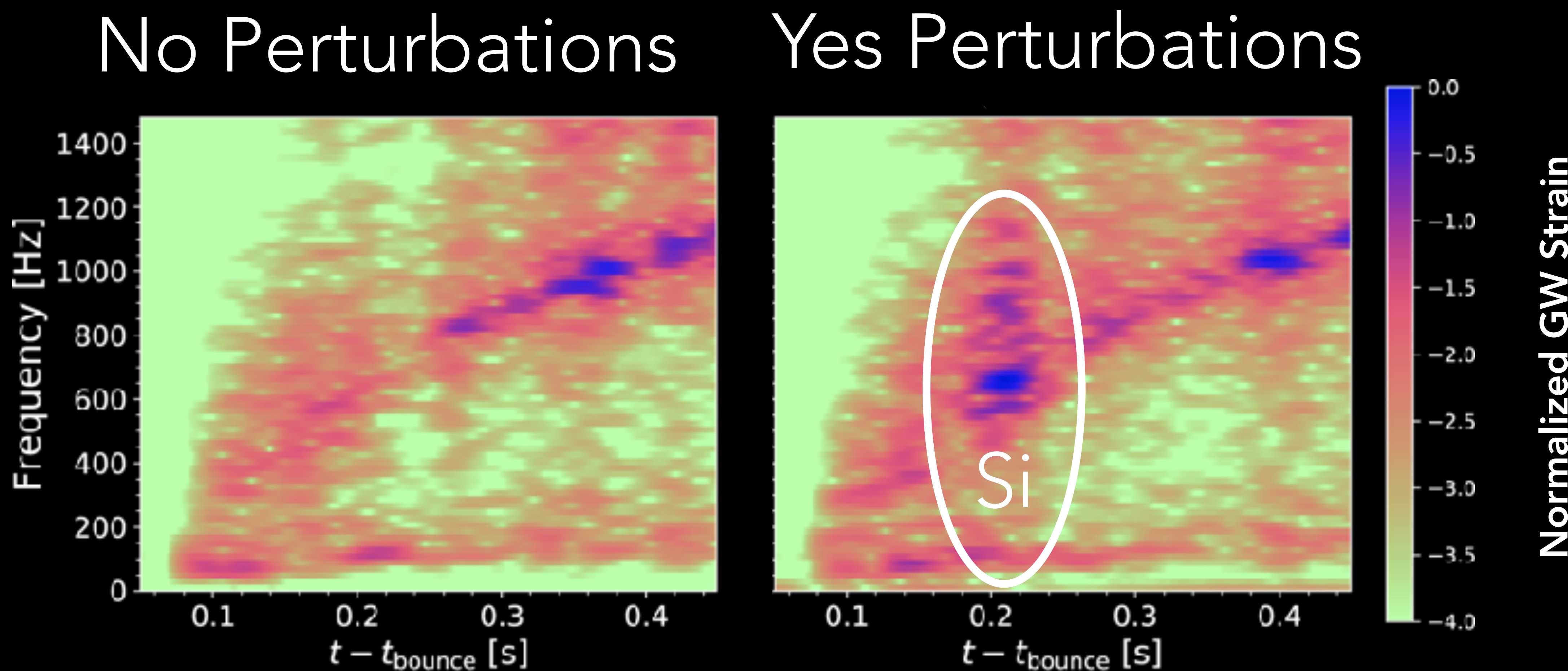
Preliminary



Slice of entropy in the x-y plane for 3D CCSN model ***(Fields + 2022b, in prep.)***.

IMPACT ON MULTI-MESSENGER ASTRONOMY

Impact of 3D progenitor on GW emission?

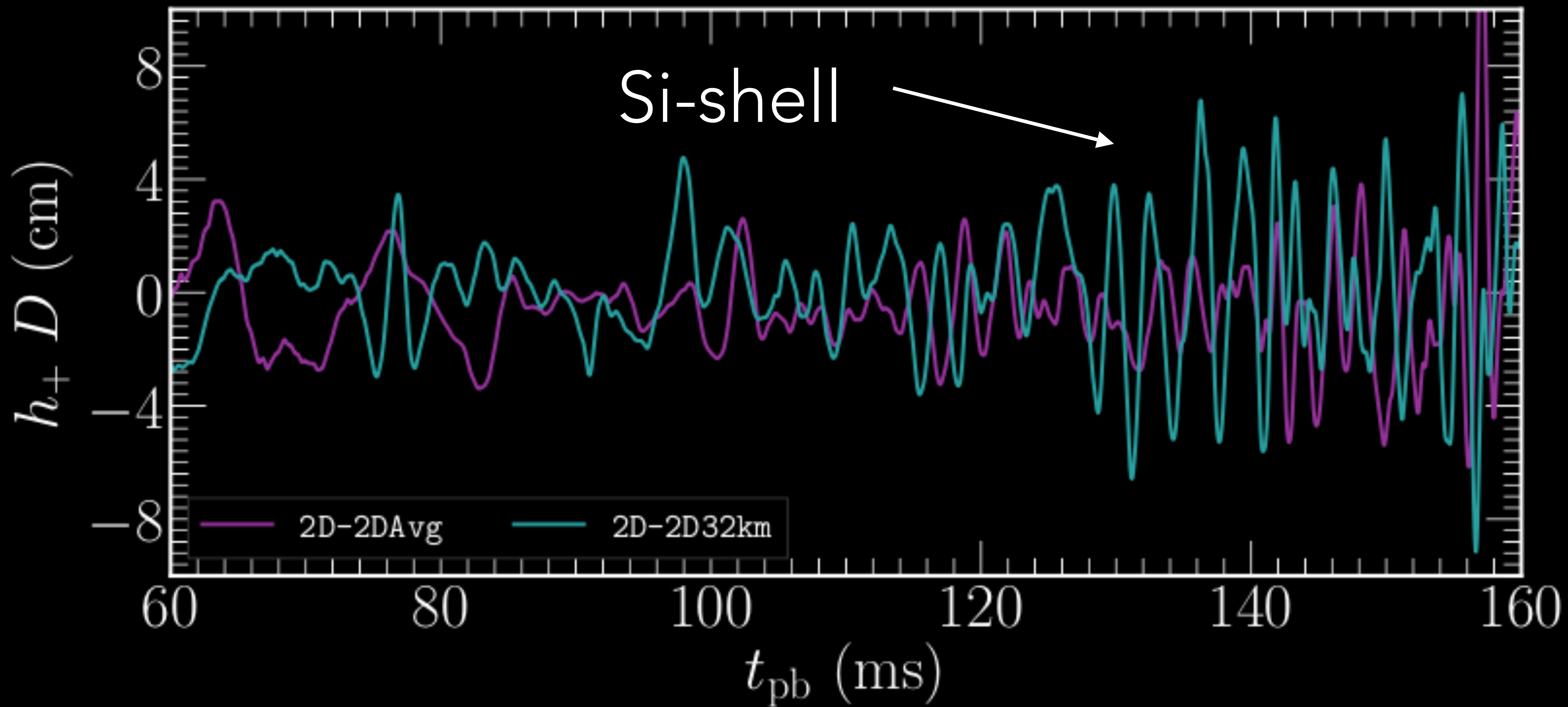


(O'Connor & Couch, 2018)

Si-shell perturbations shown in GW emission.

CCSN EXPLOSIONS OF MULTI-D PROGENITORS

Impact of perturbations on GW emission?

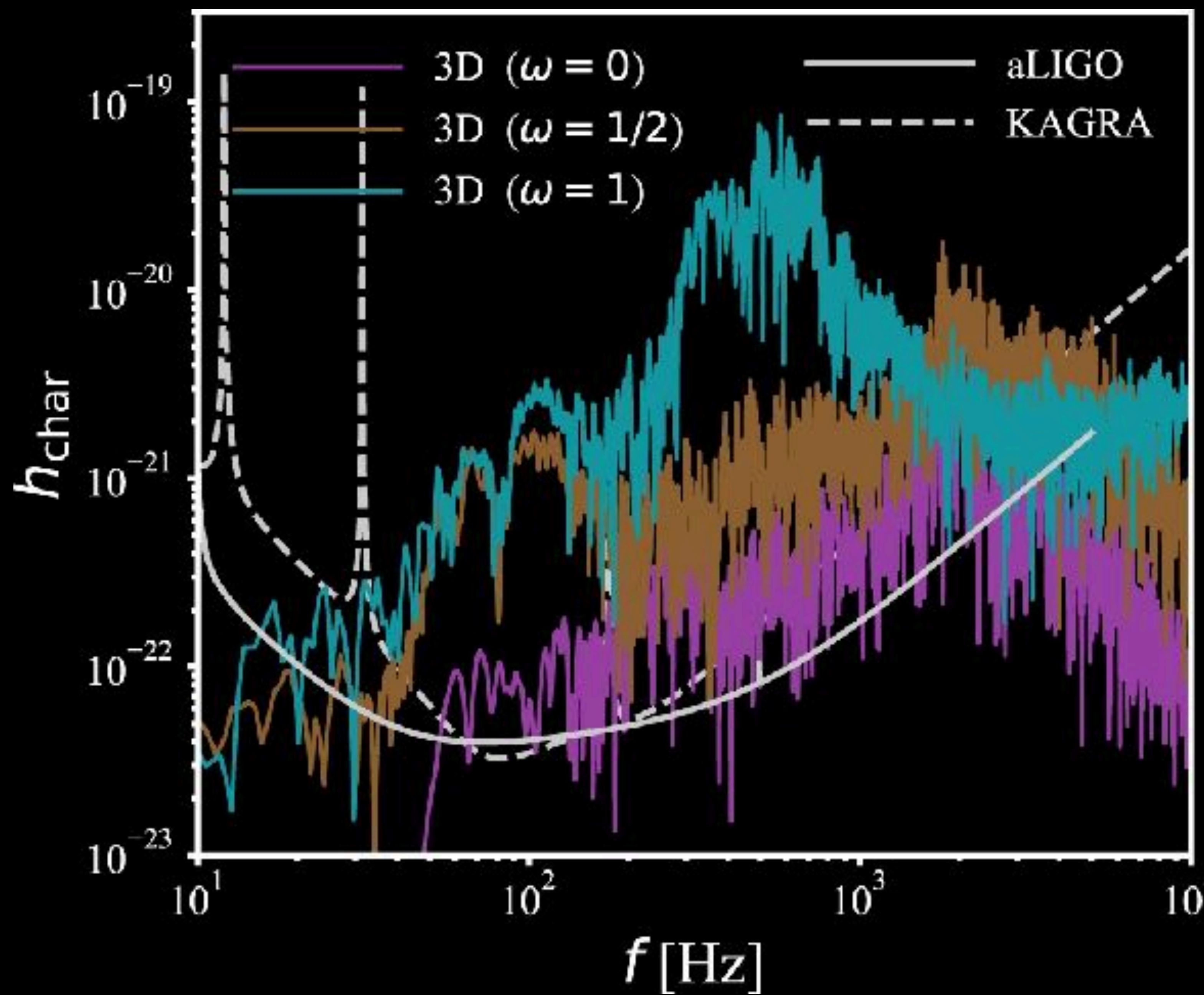


(Fields + 2022b, in prep.).

Si-shell perturbations shown in GW for $f_{\rm GW} \sim 150 - 600$ (Hz).

CCSN EXPLOSIONS OF MULTI-D PROGENITORS

- 3D rotating explosion models.
- Detectable at 10 kpc (60 kpc w/LEN).
- Rotation can amplify signal.



(Pan+ 2021, ApJ).

CONCLUSIONS & SUMMARY

3D models of stellar convection necessary for accurate description of state of model near collapse

(Fields & Couch, 2020, ApJ; Fields & Couch 2021, ApJ)

- Convection occurring at many scales, large dominant mode near collapse
- 3D instabilities can affect flow properties and mass entrainment
- Mach number profiles show favorable conditions for explosion.

3D rotating progenitor models ALSO necessary

(Fields, 2022 arXiv:2112.12800)

- Redistribution of AM diverges from MESA model. Implications for remnant.
- Turbulent transport of AM in convective shell regions.

Multi-D models can provide input for successful CCSN models

(Fields, 2022b, in prep.)

- Larger non-radial kinetic energy when using multi-D progenitor input
- 3D CCSN model showed prompt convection, asymmetric shock runaway
- Explosion properties suggest robust impact on multi-messenger signals

LOOKING FORWARD

Magnetic Fields

- Field amplification in pre-supernova phase and collapse
- Long term strength and topology during explosion

Neutrinos

- Low energy neutrinos during pre-supernova phase - impact of 3D structure
- Coupling neutrino emission properties with GW signals of 3D explosions

Angular momentum transport

- 3D redistribution affecting our compact object estimates
- Feedback into AM transport assumed in 1D models

THANK YOU

Questions?

Our data are online and available publicly!

doi.org/10.5281/zenodo.3976246

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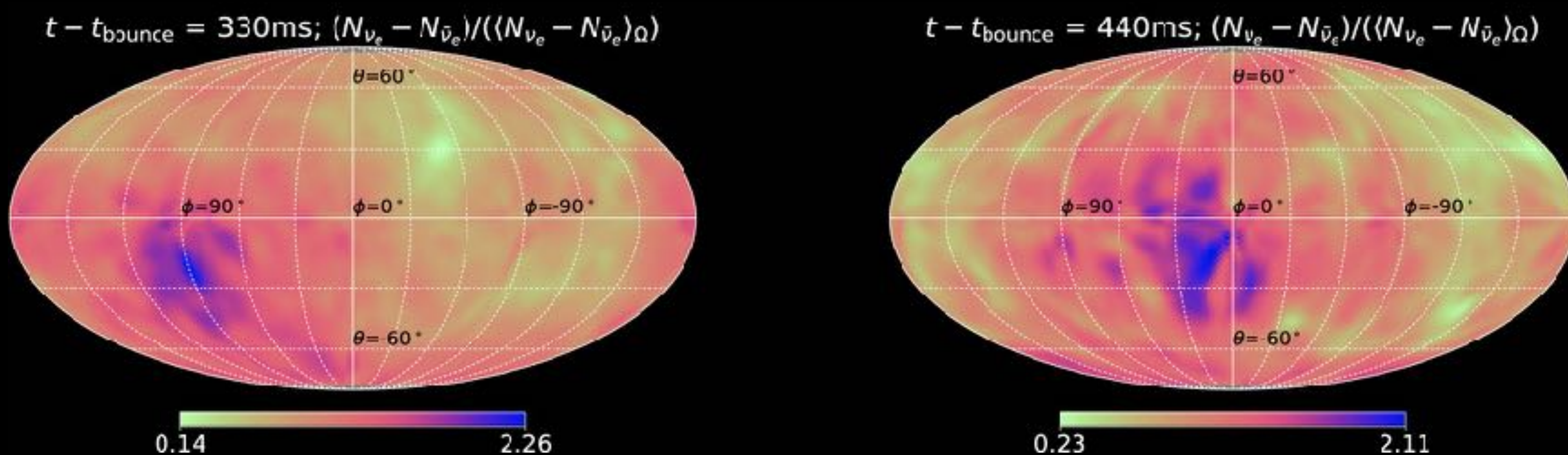
IMPACT ON MULTI-MESSENGER ASTRONOMY

Impact of 3D progenitor on neutrino emission?

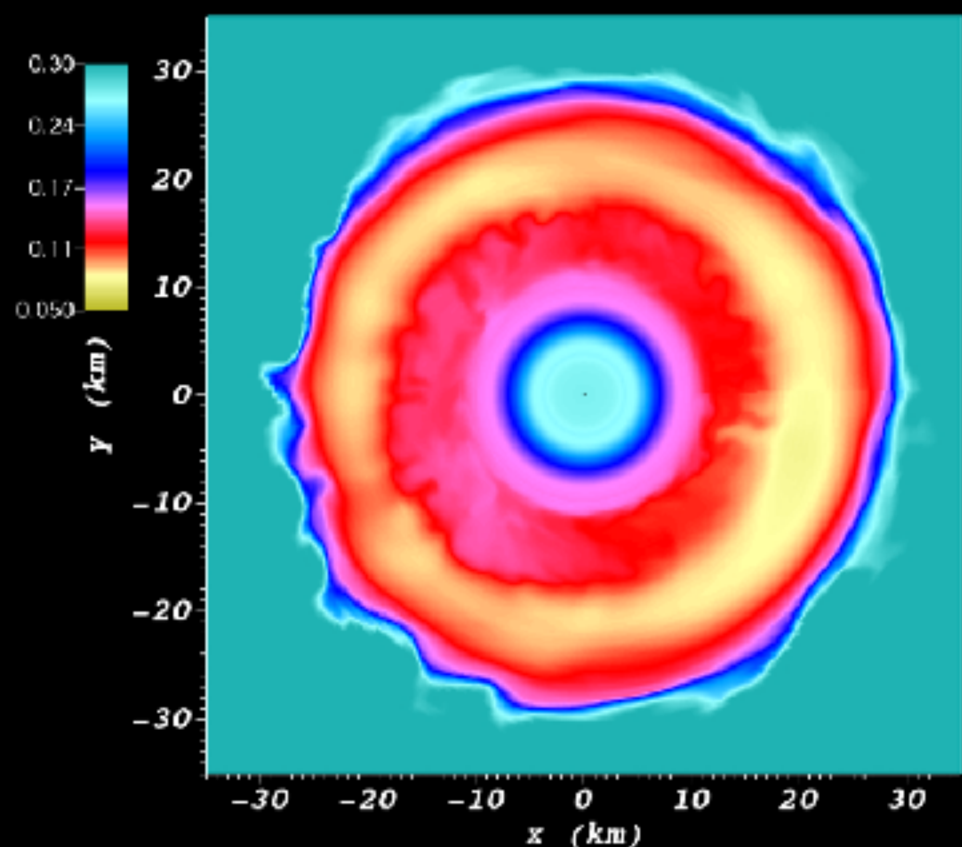


(O'Connor & Couch, 2018)

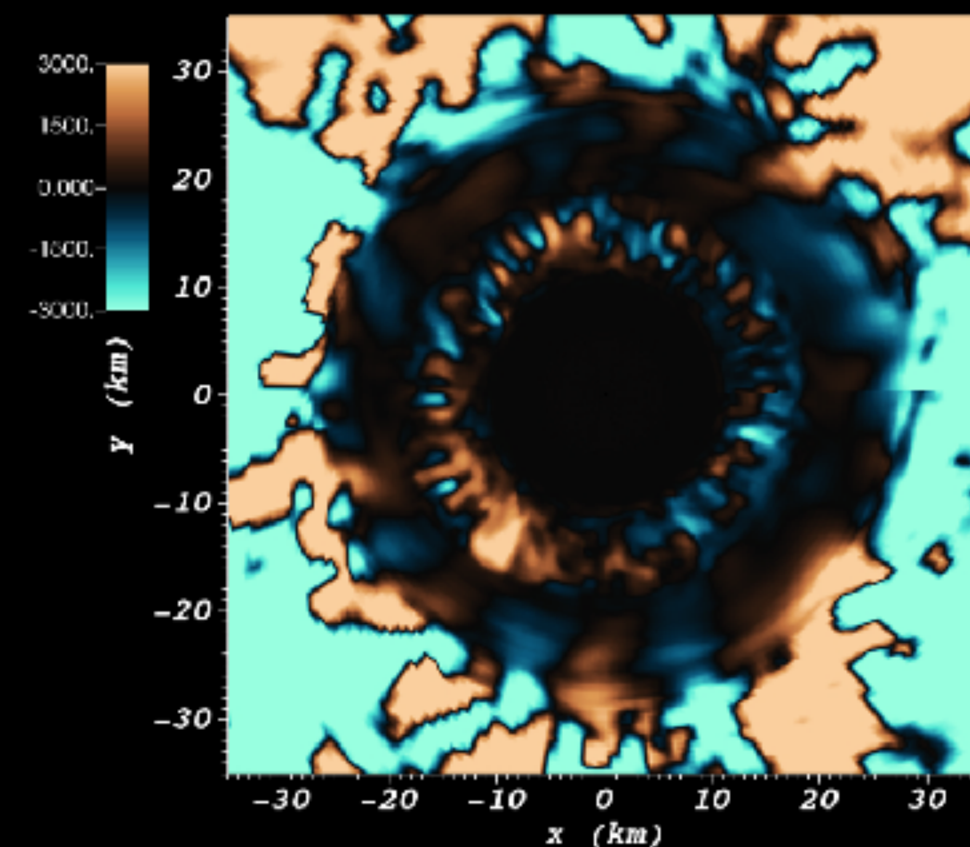
lepton-number emission self- sustained asymmetry
- **LESA** found in 3D CCSN model.

IMPACT ON MULTI-MESSENGER ASTRONOMY

Y_e



v_{radial}



(Muller+, 2020)

$$M_{\text{ZAMS}} = 18 M_{\odot} \qquad t_{\text{pb}} \sim 453 \text{ (ms)}$$

Asymmetry in electron fraction, not in radial velocity - signature of **LESA.**

IMPACT ON MULTI-MESSENGER ASTRONOMY

MNRAS 000, 1–21 (2021) Preprint 27 September 2021 Compiled using MNRAS LaTeX style file v3.0

The Collapse and Three-Dimensional Explosion of Three-Dimensional Massive-star Supernova Progenitor Models

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(arxiv.org/abs/2109.10920)

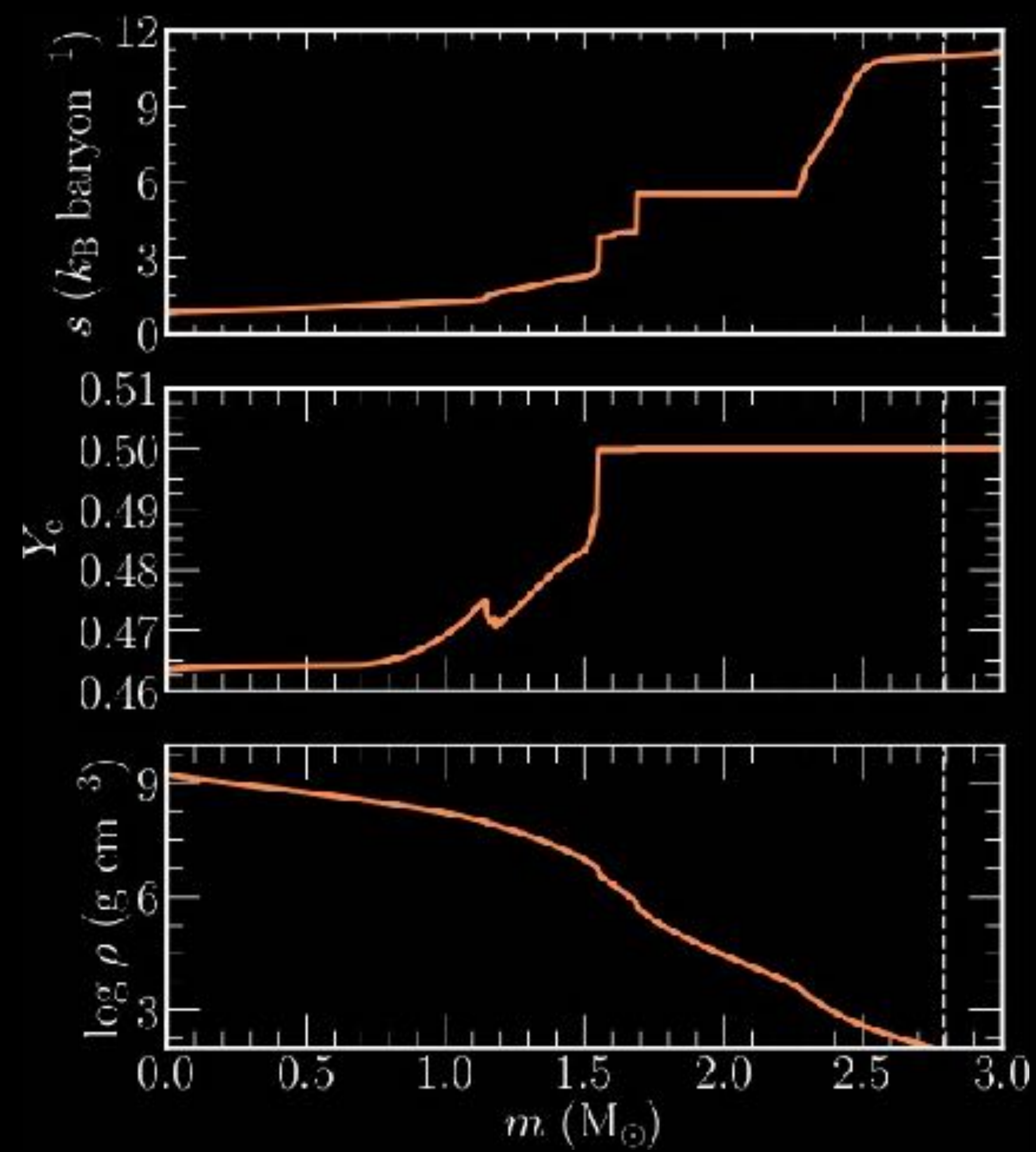
Other groups using 3D progenitors as input. Check out this recent work!

3D CCSN PROGENITORS

3D Simulations of a 15 $M_\odot$ star

MULTI-DIMENSIONAL SIMULATIONS OF MASSIVE STARS

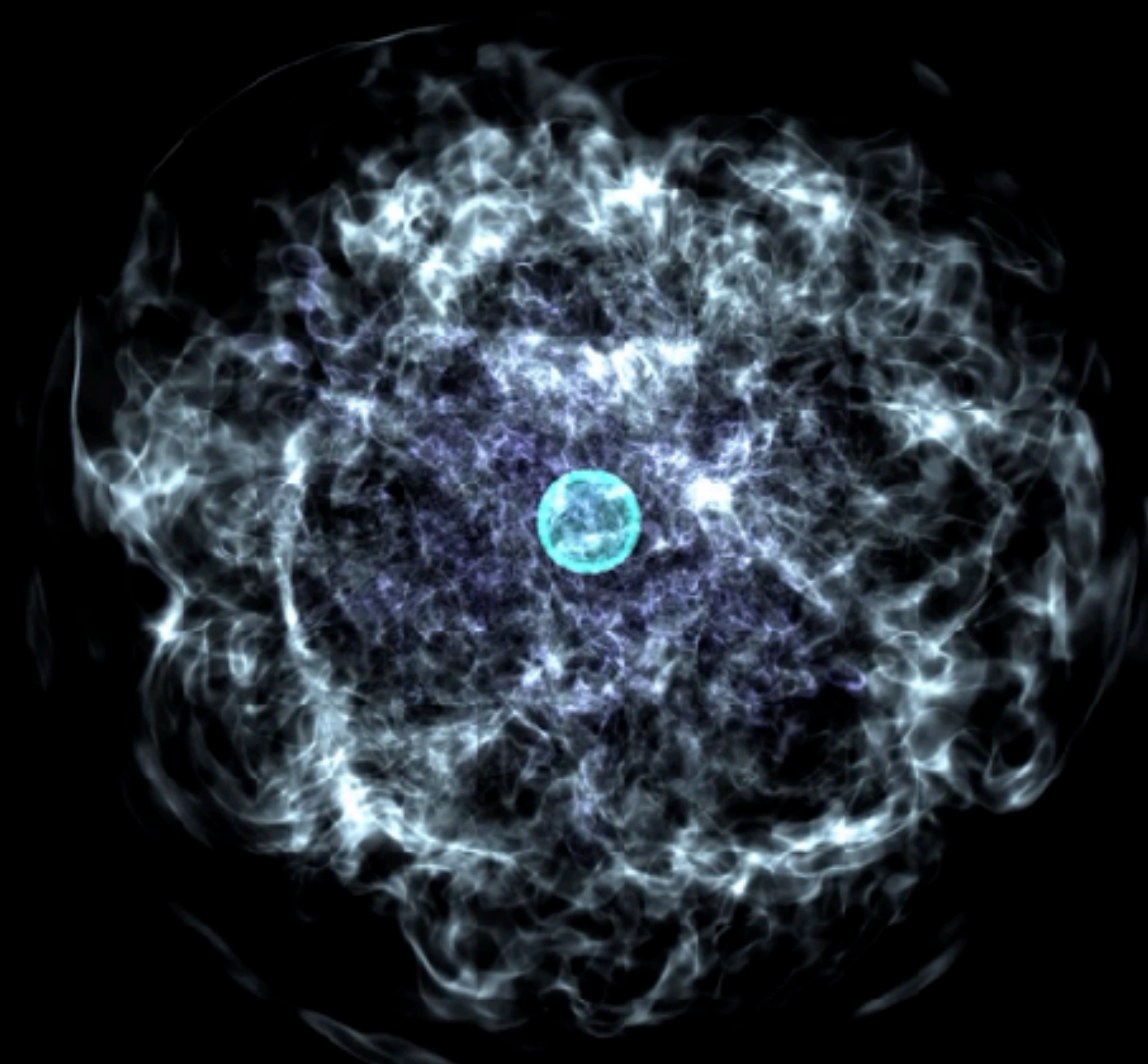
- 2/3D Hydrodynamic simulations using FLASH.
- Evolved ~**7 minutes** collapse using approximate network.
- 15 $M_\odot$ progenitor.



Stellar input model profiles from *Fields & Couch 2020.*

MULTI-DIMENSIONAL SIMULATIONS OF MASSIVE STARS

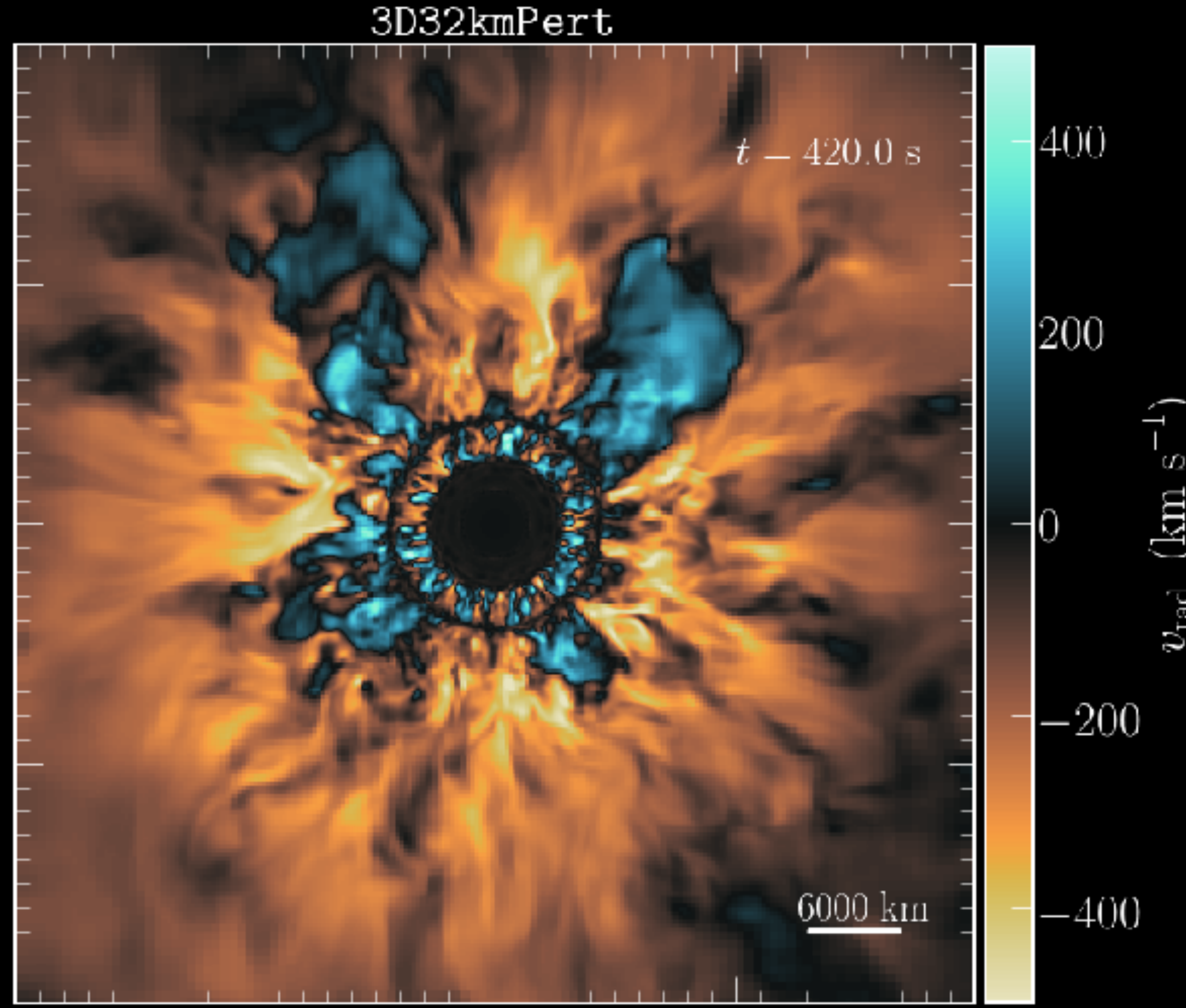
- 3D model evolved using FLASH.
- Shell convection occurring at many scales.
- Perturbations imply indirect **increase** in effective neutrino heating efficiency.



Volume rendering of the velocity field for 3D progenitor model near collapse ***(Fields & Couch 2020)***.

MULTI-DIMENSIONAL SIMULATIONS OF MASSIVE STARS

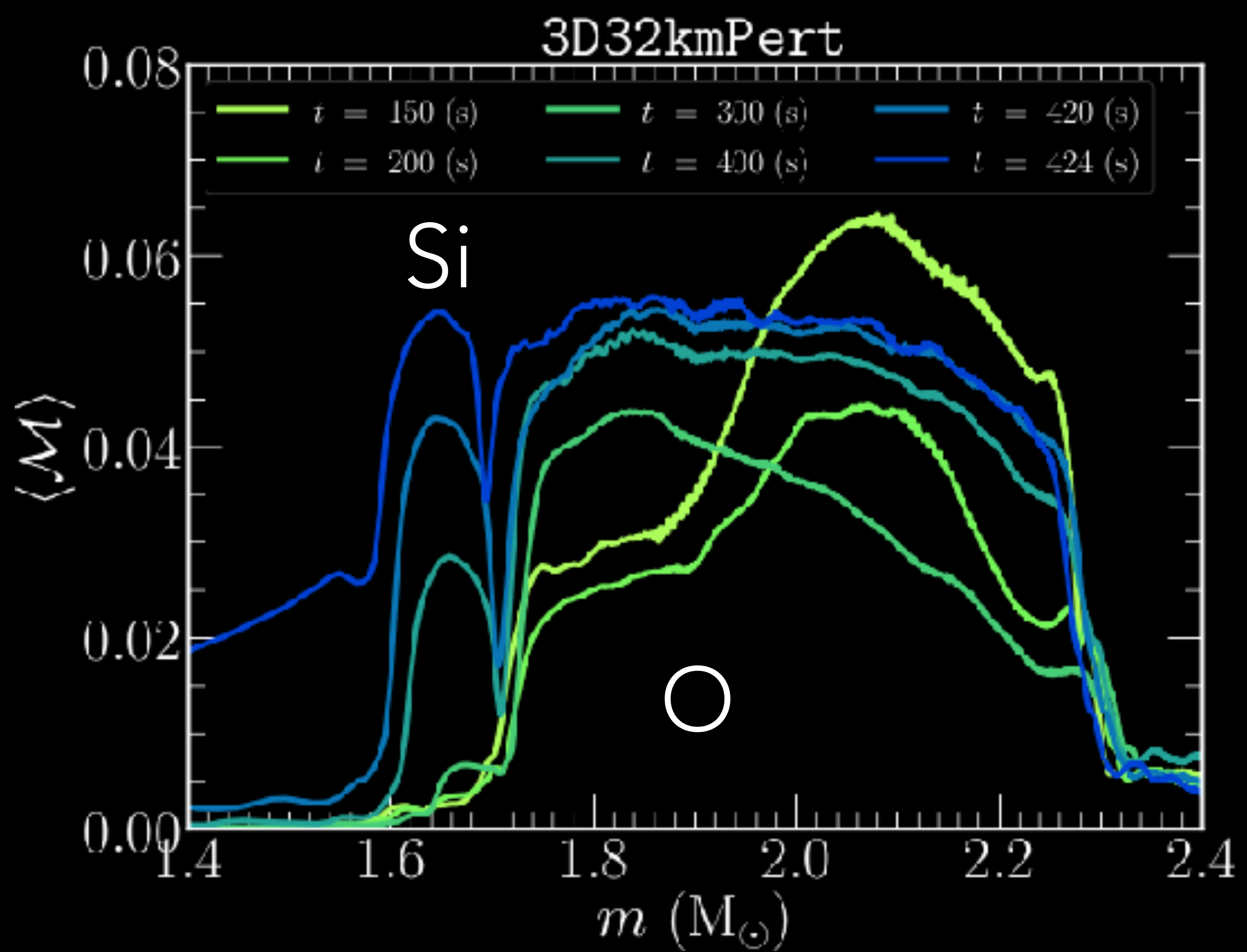
- 4 pi 3D model shows large scale plumes.
- Strong Si-shell convection.
- Convective speeds of several hundred km/s.



Slice of the radial velocity field of 3D progenitor model a few seconds before collapse ***(Fields & Couch 2020)***.

MULTI-DIMENSIONAL SIMULATIONS OF MASSIVE STARS

- Significant increase in Si-shell mach numbers at late time.
- Oxygen-shell reaches steady values early on.
- Values in O-shell lower than previous studies (Muller+2016)



Angle average mach number profiles for 3D model at different times ***(Fields & Couch 2020).***